
Planck 2015 results

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XV. Gravitational lensing


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ABSTRACT

We present the most significant measurement of the cosmic microwave background (CMB) lensing potential to date (at a level of 40σ), using temperature and polarization data from the Planck 2015 full-mission release. Using a polarization-only estimator, we detect lensing at a significance of 5σ. We cross-check the accuracy of our measurement using the wide frequency coverage and complementarity of the temperature and polarization measurements. Public products based on this measurement include an estimate of the lensing potential over approximately 70% of the sky, an estimate of the lensing potential power spectrum in bandpowers for the multipole range 40 ≤ ℓ ≤ 400, and an associated likelihood for cosmological parameter constraints. We find good agreement between our measurement of the lensing potential power spectrum and that found in the ACDM model that best fits the Planck temperature and polarization power spectra. Using the lensing likelihood alone we obtain a percent-level measurement of the parameter combination σ8Ωm0.25 = 0.591 ± 0.021. We combine our determination of the lensing potential with the E-mode polarization, also measured by Planck, to generate an estimate of the lensing B-mode. We show that this lensing B-mode estimate is correlated with the B-modes observed directly by Planck at the expected level and with a statistical significance of 10σ, confirming Planck’s sensitivity to this known sky signal. We also correlate our lensing potential estimate with the large-scale temperature anisotropies, detecting a cross-correlation at the 3σ level, as expected because of dark energy in the concordance ACDM model.

Key words. gravitational lensing: weak; cosmological parameters; cosmic microwave background: large-scale structure of Universe; cosmology: observations

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A15, page 1 of 28
1. Introduction

The cosmic microwave background (CMB) gives us a direct measurement of the early Universe when it first became transparent to radiation just 375 000 years after the Big Bang. It contains distinct signatures of the later Universe as well, which were imprinted by the process of gravitational lensing. The CMB photons that we observe today last scattered approximately 14 billion years ago, and have travelled most of the way across the observable Universe to reach us. During their journey, their paths were distorted by the gravitational tug of intervening matter, a subtle effect that may be measured statistically with high angular resolution, low-noise observations of the CMB, such as those provided by Planck. In this paper, we present lensing measurements that are based on the full-mission 2015 data release. We produce the most powerful measurement of CMB lensing to date with a 2.5% constraint on the amplitude of the lensing potential power spectrum (or alternatively, a 40$\sigma$ detection of lensing effects).

The effect of lensing is to remap the CMB fluctuations, so that the observed anisotropy in direction $\hat{n}$ is in fact the un lensed, “primordial” anisotropy in the direction $\hat{h} + \nabla \phi(\hat{h})$, where $\phi(\hat{h})$ is the CMB lensing potential defined by (e.g. Lewis & Challinor 2006)

$$\phi(\hat{h}) = -2 \int_0^\chi \frac{d\chi}{f_k(\chi)} \frac{f_k(\chi_\eta) - f_k(\chi)}{f_k(\chi_\eta) f_k(\chi)} \Psi(\chi; \eta_0 - \chi). \quad (1)$$

Here, $\chi$ is conformal distance (with $\chi_\eta \approx 14 000 \text{ Mpc}$ denoting the distance to the CMB last-scattering surface). The angular-diameter distance $f_k(\chi)$ depends on the curvature of the Universe, and is given by

$$f_k(\chi) = \begin{cases} K^{-1/2} \sin(K^{1/2}\chi) & \text{for } K > 0 \text{ (closed)}, \\ \chi & \text{for } K = 0 \text{ (flat)}, \\ |K|^{-1/2} \sinh(|K|^{1/2}\chi) & \text{for } K < 0 \text{ (open)}. \end{cases} \quad (2)$$

Finally, $\Psi(\chi; \eta)$ is the (Weyl) gravitational potential at conformal distance $\chi$ along the direction $\hat{h}$ at conformal time $\eta$ (the conformal time today is denoted as $\eta_0$). The lensing potential is an integrated measure of the mass distribution back to the last-scattering surface. The power spectrum of the lensing potential probes the matter power spectrum, which is sensitive to “late-time” parameters that modify the growth of structure such as neutrino mass (Smith et al. 2009). The amplitude of lensing effects is also a sensitive probe of geometrical parameters, such as the curvature of the Universe.

The lens-induced remapping imparts distinctive statistical signatures onto the observed CMB fluctuations, which can be mined for cosmological information in a process known as lens reconstruction (Okamoto & Hu 2003). The past several years have been seen dramatic improvements in CMB lensing measurements, moving from first detections in cross-correlation (Smith et al. 2007; Hirata et al. 2008) to cosmologically useful measurements of the lensing potential power spectrum (Das et al. 2011; van Engelen et al. 2012; Planck Collaboration XVII 2014). Recently, ground-based experiments have been able to detect the effects of lensing in polarization data as well (Hanson et al. 2013; POLARBEAR Collaboration 2014a,b; van Engelen et al. 2015).

The results in this paper extend our earlier results in Planck Collaboration XVII (2014). These were based on the March 2013 Planck$^1$ nominal-mission data release, which contains approximately 15 months of temperature data alone. The additional information in the full-mission dataset, with 30 months of Planck HFI temperature and polarization data (Planck Collaboration VII 2016; Planck Collaboration VIII 2016), allows us to improve our reconstruction noise levels by roughly a factor of two. Approximately half of this improvement comes from lower noise levels in temperature, and the other half from inclusion of polarization data. The improved lensing map is included as part of the Planck 2015 public data release (Planck Collaboration I 2016), as well as an estimate of the lensing potential power spectrum and associated likelihoods. In this paper we describe the creation of these products, as well as first science results based on them. We highlight the following science results.

- We detect lensing $B$-modes in the Planck data at a significance of $10\sigma$, using both a cross-correlation approach with the cosmic infrared background (CIB) as a tracer of the lensing potential, as well as a CMB-only approach using the TTEB trispectrum. This provides an important confirmation that Planck is sensitive to this known source of $B$-modes on intermediate and small scales.
- We make an improved measurement of the 3-point function (bispectrum) that is induced in the CMB by the correlation between lensing and the integrated Sachs-Wolfe (ISW) effect, the latter being sourced by late-time acceleration. The lensing-ISW bispectrum is now detected at the $3\sigma$ level.
- Using only lensing information (along with well-motivated priors), we constrain the parameter combination $\sigma_8\Omega_m^{25}$ to roughly 3%.

The organization of this paper is as follows. In Sect. 2 we give a summary of our analysis pipeline for the full-mission data. Our analysis is very similar to that presented in Planck Collaboration XVII (2014), with straightforward extensions to polarization data, as well as a few additional improvements, and so we have deferred most technical details to appendices. In Sect. 3 we present our main results, including the lensing potential map, bandpower estimates of the lensing potential power spectrum, and implications for cosmological parameters. In Sect. 4 we present a suite of consistency and null tests to verify our bandpower estimates, and in Sect. 5 we conclude. A series of appendices provide more technical details of our analysis pipeline, modeling of the CIB, and dependence of the lensing power spectrum on $\Lambda$CDM parameters.

2. Data and methodology

Here we give a brief overview of the procedure that we use to measure the CMB lensing potential and its power spectrum from the Planck maps. We defer the detailed technical aspects of our pipeline to Appendix A.

Our main results are based on a foreground-cleaned map of the CMB synthesized from the raw Planck 2015 full-mission frequency maps using the SMICA code (Planck Collaboration IX 2016). This foreground-cleaned map combines all nine frequency bands from 30 GHz to 857 GHz with scale-dependent
coefficients chosen to provide unit response to the CMB with minimal variance; for more details see Planck Collaboration IX (2016)\(^2\). On the small scales where lensing effects are most important (multipoles above \(\ell = 1000\), most of the CMB information in the foreground-cleaned maps originates with the 143 GHz and 217 GHz channel data. These channels have beams that are approximately Gaussian with full-width-at-half-maximum (FWHM) parameters of 7\(^\circ\) and 5\(^\circ\), respectively. Their noise is approximately white for 1000 \(\leq \ell \leq 3000\), with levels of 30 \(\mu K\) arcmin (60 \(\mu K\) arcmin) in temperature (polarization) at 143 GHz, and 40 \(\mu K\) arcmin (95 \(\mu K\) arcmin) at 217 GHz.

We reconstruct the CMB lensing potential using quadratic estimators that exploit the statistical anisotropy induced by lensing, following Okamoto & Hu (2003). Neglecting the lensing of primordial \(B\)-modes there are five possible estimators, denoted by \(\hat{\phi}^{TT}\), \(\hat{\phi}^{TE}\), \(\hat{\phi}^{EE}\), \(\hat{\phi}^{EB}\), and \(\hat{\phi}^{TB}\) which are based on various correlations of the CMB temperature \((T)\) and polarization \((E\) and \(B\)). In addition, we can form a minimum-variance estimator that combines all five estimators, which we denote as \(\hat{\phi}^{MV}\). In Fig. 1 we plot the lens reconstruction noise levels for these estimators. The most powerful estimator is \(TT\), although the \(TE\) and \(EE\) estimators are also useful on large angular scales.

The quadratic lensing estimators take inverse-variance filtered CMB multipoles as input. We obtain these using a filter that masks the Galaxy and point sources, and also bandpass filters the data in harmonic space to 100 \(\leq \ell \leq 2048\). For our baseline analysis, we start with the masks used in the analysis of Planck Collaboration XVII (2014), with a Galaxy mask that removes 30.2% of the sky and a point-source mask that removes an additional 0.7% of the sky. This removes the brightest Sunyaev-Zel’dovich (SZ) clusters as described in Planck Collaboration XVII (2014); contamination of the lensing reconstruction from residual unresolved SZ is expected to be at the percent level for Planck (van Engelen et al. 2014), even without component

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3. We use masks with \(S/N \geq 5\) in the Planck cluster catalogue that accompanied the 2013 release (PSZ1; Planck Collaboration XXIX 2014). The mass limit is redshift dependent, with a high-redshift limit at 80% completeness of \(M_{500} \approx 6 \times 10^{14} M_\odot\). If we conservatively adopt this mass limit (we certainly mask low-redshift clusters to lower masses), the results of van Engelen et al. (2014) suggest that the SZ trispectrum at 143 GHz gives a positive bias of at most a few percent in the reconstructed lensing power spectrum for multipoles \(\ell < 2000\), while the lensing-SZ-SZ bispectrum gives a negative bias of similar size. The SZ signal is much smaller in the 217 GHz temperature data, which is around the null of the (thermal) SZ effect.

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4. In this paper, we use multipole indices \(LM\) for the lens reconstruction, reserving multipole indices \(lm\) for the CMB fields.
We account for this missing power simply by adding coloured Gaussian noise to the simulations to make their $TT$, $EE$, and $BB$ power spectra agree with the data. This approach implicitly assumes that any non-Gaussianity of these residual components does not couple significantly to our lensing estimates. We perform consistency tests in Sect. 4 to check the validity of these assumptions.

Throughout this paper we use a spatially-flat fiducial cosmology with baryon density given by $\omega_\text{b} = \Omega_\text{b} h^2 = 0.0222$, cold dark matter (CDM) density $\omega_\text{c} = \Omega_\text{c} h^2 = 0.1203$, neutrino energy density $\omega_\nu = \Omega_\nu h^2 = 0.00064$ (corresponding to two massless neutrinos and one massive with mass 0.06 eV), Hubble constant $H_0 = 100 h$ km s$^{-1}$ Mpc$^{-1}$ with $h = 0.6712$, spectral index of the power spectrum of the primordial curvature perturbation $n_s = 0.96$, amplitude of the primordial power spectrum (at $k = 0.05$ Mpc$^{-1}$) $A_s = 2.09 \times 10^{-9}$, and Thomson optical depth through reionization $\tau = 0.065$. These cosmological parameters also form the basis for the FFP8 Monte Carlo simulation set. In addition to rescaling the FFP8 maps as already discussed, we have also adjusted the power spectra of the fiducial model by rescaling the CMB temperature and polarization spectra by a factor of 1.0134, and the temperature-lensing cross-correlation $C_{\ell}^{T\phi}$ by 1.0134. We have not applied any scaling to the fiducial lensing power spectrum. Our reconstruction methodology (in particular, the renormalization corrections, addition of foreground power, and realization-dependent bias corrections that we apply, discussed in Appendix C) renders the cosmological interpretation of our lensing estimates insensitive to errors in the fiducial model power spectra and simulations.

3. Results

In this section, we provide a summary of the first science results obtained with the minimum-variance lens reconstruction from the Planck full-mission data. The lensing potential map is presented in Sect. 3.1, and this is combined in Sect. 3.2 with the $E$-mode polarization measured by Planck to obtain a map of the expected $B$-mode polarization due to lensing. We further show that this is correlated with the $B$-modes measured by Planck at the expected level. In Sect. 3.3, we cross-correlate the reconstructed lensing potential with the large-angle temperature anisotropies to measure the $C_{\ell}^{T\phi}$ correlation sourced by the ISW effect. Finally, the power spectrum of the lensing potential is presented in Sect. 3.4. We use the associated likelihood alone, and in combination with that constructed from the Planck temperature and polarization power spectra (Planck Collaboration XI 2016), to constrain cosmological parameters in Sect. 3.5.

3.1. Lensing potential

In Fig. 2 we plot the Wiener-filtered minimum-variance lensing estimate, given by

$$\hat{\phi}_{LM}^{\text{WF}} = \frac{C_{\ell}^{\phi\phi, \text{fid}}}{C_{\ell}^{\phi\phi, \text{fid}} + N_{\ell}^{\phi\phi}} \hat{\phi}_{LM}^{\text{MV}},$$

where $C_{\ell}^{\phi\phi, \text{fid}}$ is the lensing potential power spectrum in our fiducial model and $N_{\ell}^{\phi\phi}$ is the noise power spectrum of the reconstruction. As we shall discuss in Sect. 4.5, the lensing potential estimate is unstable for $L < 8$, and so we have excluded those modes for all analyses in this paper, as well as in the MV lensing map.

Fig. 2. Lensing potential estimated from the SKICA full-mission CMB maps using the MV estimator. The power spectrum of this map forms the basis of our lensing likelihood. The estimate has been Wiener filtered following Eq. (5), and band-limited to $8 \leq L \leq 2048$.

Fig. 3. Simulation of a Wiener-filtered MV lensing reconstruction (upper) and the input $\phi$ realization (lower), filtered in the same way as the MV lensing estimate. The reconstruction and input are clearly correlated, although the reconstruction has considerable additional power due to noise.

As a visual illustration of the signal-to-noise level in the lensing potential estimate, in Fig. 3 we plot a simulation of the MV reconstruction, as well as the input $\phi$ realization used. The reconstruction and input are clearly correlated, although the reconstruction has considerable additional power due to noise. As can be seen in Fig. 1, even the MV reconstruction only has $S/N \approx 1$ for a few modes around $L \approx 50$.

The MV lensing estimate in Fig. 2 forms the basis for a public lensing map that we provide to the community (Planck Collaboration I 2016). The raw lensing potential estimate has a very red power spectrum, with most of its power on large angular scales. This can cause leakage issues when cutting the map (for example to cross-correlate with an additional mass tracer over a small portion of the sky). The lensing convergence $\kappa$ defined by

$$\kappa_{LM} = \frac{L(L+1)}{2} \phi_{LM},$$

A15, page 4 of 28
has a much whiter power spectrum, particularly on large angular scales. The reconstruction noise on $\kappa$ is approximately white as well (Bucher et al. 2012). For this reason, we provide a map of the estimated lensing convergence $\kappa$ rather than the lensing potential $\phi$.

### 3.2. Lensing B-mode power spectrum

The odd-parity $B$-mode component of the CMB polarization is of great importance for early-universe cosmology. At first order in perturbation theory it is not sourced by the scalar fluctuations that dominate the temperature and polarization anisotropies, and so the observation of primordial $B$-modes can be used as a uniquely powerful probe of tensor (gravitational wave) or vector perturbations in the early Universe. A detection of $B$-mode fluctuations on degree angular scales, where the signal from gravitational waves is expected to peak, has recently been reported at 150 GHz by the BICEP2 collaboration (Ade et al. 2014). Following the joint analysis of BICEP2 and Keck Array data (also at 150 GHz) and the Planck polarization data, primarily at 353 GHz (BICEP2/Keck Array and Planck Collaborations 2015), it is now understood that the $B$-mode signal detected by BICEP2 is dominated by Galactic dust emission. The joint analysis gives no statistically-significant evidence for primordial $B$-modes, and establishes a 95% upper limit $r_{0.05} < 0.12$. This still represents an important milestone for $B$-mode measurements, since the direct constraint from the $B$-mode power spectrum is now as constraining as indirect, and model-dependent, constraints from the $TT$ spectrum (Planck Collaboration XIII 2016).

In addition to primordial sources, the effect of gravitational lensing also generates $B$-mode polarization. The displacement of lensing maps $E$-mode polarization into $B$-mode as (Smith et al. 2009)

$$B^{\text{lens}}_{\ell_{\text{Emg}}} = (-1)^{m_B} \sum_{LM} \sum_{\ell \equiv m_B} \left( \frac{\ell_E}{m_E} \frac{\ell_B}{m_B} \frac{L}{M} \right) W^{\phi_E}_{\ell_E m_E} E_{\ell m_B} \phi_{LM},$$

where $W^{\phi_E}_{\ell_E m_E}$ is a weight function and the bracketed term is a Wigner-3j symbol. On scales $\ell_B \approx 1000$ the lensing $B$-mode power spectrum resembles that of white noise, with a level of about $5 \mu K$ arcmin. This lensing power acts as a potential source of confusion for the measurement of primordial $B$-modes, which can be estimated and ultimately removed in a process of delensing. Given an estimate for the lensing potential $\phi$ and the $E$-mode polarization measured by Planck we can synthesize a lensed $B$-mode map for this purpose using Eq. (7). The $5 \mu K$ arcmin level of the lensing $B$-mode power spectrum is an order of magnitude lower than the Planck 2015 noise levels, and so delensing does not significantly improve our $B$-mode measurements; however, the cross-correlation of the lensing $B$-mode template with the observed $B$-mode sky provides a useful check on the ability of Planck to measure this known source of $B$-modes.

We show the results of such a cross-correlation in Fig. 4, finding good agreement with the expected lensing $B$-mode power spectrum. In addition to our fiducial MV lensing potential estimate, which uses both temperature and polarization data, we have also estimated the lensing $B$-mode power spectrum using the $TT$-only lensing estimator to measure $\phi$, as well as the CIB fluctuations measured by the 545 GHz Planck channel.

$$\tilde{A}^B_{8-2048} = 0.93 \pm 0.08 \quad (\phi^{\text{MV}}),$$
$$\tilde{A}^B_{8-2048} = 0.95 \pm 0.09 \quad (\phi^{TT}),$$
$$\tilde{A}^B_{8-2048} = 0.93 \pm 0.10 \quad (\text{CIB})$$

for the three estimates, each corresponding to a roughly 10$\sigma$ detection of lensing $B$-mode power in the Planck data. The shape of the cross-correlation is also in good agreement with expectation. Taking the bins in Fig. 4 as independent, forming a $\chi^2$ relative to the theory model and comparing to the distribution from simulations we obtain probability-to-exceed (PTE) values of 48%, 71%, and 78% using the MV, $TT$, and CIB lensing estimates, respectively.

### 3.3. Lensing-ISW bispectrum

As photons travel towards us from the last scattering surface, they are not only deflected by gravitational lensing, they also receive net red/blue shifts from gravitational potentials that evolve if they are crossed at late times. This phenomenon, known as the integrated Sachs-Wolfe (ISW) effect, is believed to generate anisotropies in the observed CMB temperature on large ($\ell \approx 100$) angular scales. It is of particular interest because the decay of gravitational potentials, which produces the ISW effect, does not occur during matter domination, but only at redshifts $z \lesssim 2$ when dark energy becomes dynamically important. The ISW effect can be detected statistically by cross-correlating the observed temperature anisotropies with a tracer of structure at these redshifts. Here we use the lensing potential, which is well-matched to the ISW effect (Hu 2002). Current Planck results

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5 To calculate the scaling between the CIB map and the lensing potential $\phi$, we model the CIB using the simple model of Hall et al. (2010) for its redshift and frequency dependence. Further details are given in Appendix D.
In Planck Collaboration XVII (2014), using the $TT$ lensing estimator on the multipole range $10 \leq L \leq 100$ we measured a somewhat higher value for the lensing-ISW bispectrum amplitude of $\hat{A}_{T^6, 2013} = 0.85 \pm 0.35$ \(^\dagger\). We expect the difference with respect to the 2015 $TT$ measurement to have standard deviation of approximately $\sqrt{0.35^2 - 0.32^2} = 0.14$, and so the observed difference of $0.85 - 0.68 = 0.17$ is reasonable.

### 3.4. Lensing potential power spectrum

In Fig. 6 we plot our estimate of the lensing potential power spectrum obtained from the MV reconstruction, as well as several earlier measurements. We see good agreement with the shape in the fiducial model, as well as earlier measurements (detailed comparisons with the 2013 spectrum are given later in this section). In Sect. 4 we perform a suite of internal consistency and null tests to check the robustness of our lensing spectrum to different analysis and data choices.

We estimate the lensing potential power spectrum in bandpowers for two sets of bins: a “conservative” set of eight uniformly-spaced bins with $\Delta L = 45$ in the range $40 \leq L \leq 400$; and an “aggressive” set of 18 bins that are uniformly spaced in $L^{10}$ over the multipole range $8 \leq L \leq 2048$. The conservative bins cover a multipole range where the estimator signal-to-noise is greatest. They were used for the Planck 2013 lensing likelihood described in Planck Collaboration XVII (2014). The aggressive bins provide good sensitivity to the shape of the lensing power on both large and small scales, however they are more easily biased by errors in the mean-field corrections (which are large at $L < 40$) and the disconnected noise bias corrections (at $L > 400$). Results for the bandpower amplitudes $\hat{A}^6_{\ell}$ defined in Eq. (4), are given in Table 1 for both sets of multipole bins.

Nearly all of the internal consistency tests that we present in Sect. 4 are passed at an acceptable level. There is, however, mild evidence for a correlated feature in the curl-mode null-test, centred around $L \approx 500$. The $L$ range covered by this feature includes $638 \leq L \leq 762$, for which the (gradient) lensing reconstruction bandpower is $3.6\sigma$ low compared to the predicted power of the fiducial model. In tests of the sensitivity of parameter constraints to lensing multipole range, described later in Sect. 3.5.4, we find shifts in some parameters of around 1$\sigma$ in going from the conservative to aggressive range, with negligible improvement in the parameter uncertainty. Around half of these shifts come from the outlier noted above. For this reason, we adopt the conservative multipole range as our baseline here, and in other Planck 2015 papers, when quoting constraints on cosmological parameters. However, where we quote constraints on amplitude parameters in this paper, we generally give these for both the aggressive and conservative binning. The aggressive bins are also used for all of the $C^6_{\ell}$ bandpower plots in this paper.

Estimating an overall lensing amplitude (following Eq. 4) relative to our fiducial theoretical model, for a single bin over both the aggressive and conservative multipole ranges we find

$$\hat{A}_{40 \rightarrow 400}^{6, \text{MV}} = 0.987 \pm 0.025,$$
$$\hat{A}_{8 \rightarrow 2048}^{6, \text{MV}} = 0.983 \pm 0.025.$$

\(^\dagger\) The amplitude quoted in Planck Collaboration XVII (2014) is actually $\hat{A}_{T^6, 2013} = 0.78 \pm 0.32$, however this is measured with respect to a slightly different fiducial cosmology than the one used here. Those measurements have been renormalized to the fiducial model used for this paper with a factor of 1.09.
These measurements of the amplitude of the lensing power spectrum both have precision of 2.5%, and are non-zero at 40/π.

Given the measured amplitude, it is clear that our overall lensing amplitude estimate is consistent with our fiducial ΛCDM model (which has $A = 1$). The shape of our measurement is also in reasonable agreement with the fiducial model. Marginalizing over an overall amplitude parameter, for the 19 bins plotted in Fig. 6 we obtain a $\chi^2$ with respect to our fiducial model of 28 (with $N_{dof} = 18$), with a corresponding PTE of 6%. A large portion of this $\chi^2$ is driven by the outlier bandpower with $638 \leq L \leq 762$. Removing this bandpower gives a $\chi^2$ of 16.5 ($N_{dof} = 17$), with a PTE of 49%. For the more conservative multipole range $40 \leq L \leq 400$, using eight linear bins we obtain $\chi^2$ of 8.9 ($N_{dof} = 7$), with a PTE of 26%.

Finally, we note that the lensing bandpowers measured here are in good agreement with the Planck 2013 results (Planck Collaboration XVII 2014). The only clear visual difference between the two measurements in Fig. 6 is for the $L = 40–65$ bin, where the 2013 bandpower is significantly higher than the current measurement. As previously discussed in Planck Collaboration XVII (2014), this bin amplitude was sensitive to the choice of foreground cleaning, and decreased when using a component-separated map for the reconstruction (as is done here), and so this difference is expected. To compare measurements in further detail, we calculate a $\chi^2$ for the difference between our 2013 and 2015 MV lensing bandpower estimates assuming a diagonal covariance as

$$\chi^2_{2013-2015} = \sum_{L_b} \frac{(\hat{C}_L^{\phi\phi}_{obs,2013} - \hat{C}_L^{\phi\phi}_{obs,2015})^2}{\text{Var}(\hat{C}_L^{\phi\phi}_{obs,2013}) - \text{Var}(\hat{C}_L^{\phi\phi}_{obs,2015})},$$

where $L_b$ indexes the bandpower bins. We approximate the variance of the differences of the bandpowers as the difference of the variances in forming this $\chi^2$. Using the conservative bins for $40 \leq L \leq 400$, which were used for the 2013 lensing likelihood, we obtain a value of $\chi^2_{2013-2015} = 10.8 (N_{dof} = 8)$, with a corresponding PTE of 22%. The 2013 result is temperature only; if we compare it directly to the 2015 $TT$ lensing bandpowers, we obtain a value of $\chi^2 = 7.2 (N_{dof} = 8, \text{PTE} = 52\%)$. We note that some care should be taken when comparing the measured lensing amplitudes between 2013 and 2015, because the fiducial spectra against which they are measured differ significantly.

In 2013, for example, we measured $\hat{A}^{\phi,2013}_{40-400} = 0.943 \pm 0.040$. However, this was measured with respect to a fiducial $C_L$ that is between 6% and 7% higher than the one used here in this multipole range. Renormalizing the 2013 measurements to determine the amplitude for the 2015 fiducial cosmology, we obtain $1.005 \pm 0.043$, which can be compared directly to the measurement of $\hat{A}^{\phi,2015}_{40-400} = 0.987 \pm 0.025$ for 2015. It is interesting to note that there is a 2% difference in power between the Planck 2013 and 2015 maps from HFI due to a change in the calibration, as well as an improved beam model (Planck Collaboration VII 2016). This shift does not couple to the lensing reconstruction; lensing is a geometric effect, and so an overall change in the calibration does not affect the lens reconstruction procedure, provided that a consistent power spectrum is used to analyse the maps, as is done here.

3.5. Likelihoods and cosmological parameter constraints

In this section we discuss cosmological implications of the Planck lensing potential power spectrum estimate. We use a Gaussian likelihood in the measured bandpowers, with corrections to account for errors in the fiducial CMB power spectra that are used to normalize and debias the lensing estimates. This procedure is described in more detail in Appendix C.
Table 1. Lensing potential power spectrum estimates from the MV lens reconstruction.

<table>
<thead>
<tr>
<th>L_{\text{min}}</th>
<th>L_{\text{max}}</th>
<th>\hat{A}^6</th>
<th>10^5[L(L+1)]\hat{C}_{\ell}^{\phi\phi}/2\pi</th>
</tr>
</thead>
<tbody>
<tr>
<td>Conservative multipole range (40 \leq L \leq 400)</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>40</td>
<td>84</td>
<td>1.01 \pm 0.05</td>
<td>1.28 \pm 0.06</td>
</tr>
<tr>
<td>85</td>
<td>129</td>
<td>1.02 \pm 0.04</td>
<td>1.00 \pm 0.04</td>
</tr>
<tr>
<td>130</td>
<td>174</td>
<td>0.93 \pm 0.05</td>
<td>0.70 \pm 0.04</td>
</tr>
<tr>
<td>175</td>
<td>219</td>
<td>0.89 \pm 0.07</td>
<td>0.52 \pm 0.04</td>
</tr>
<tr>
<td>220</td>
<td>264</td>
<td>0.86 \pm 0.10</td>
<td>0.41 \pm 0.05</td>
</tr>
<tr>
<td>265</td>
<td>309</td>
<td>1.03 \pm 0.11</td>
<td>0.41 \pm 0.05</td>
</tr>
<tr>
<td>310</td>
<td>354</td>
<td>1.24 \pm 0.13</td>
<td>0.41 \pm 0.04</td>
</tr>
<tr>
<td>355</td>
<td>400</td>
<td>0.97 \pm 0.16</td>
<td>0.27 \pm 0.05</td>
</tr>
</tbody>
</table>

Aggressive multipole range (8 \leq L \leq 2048)

| 8 | 20 | 1.09 \pm 0.17 | 1.29 \pm 0.21 |
| 21 | 39 | 1.01 \pm 0.09 | 1.39 \pm 0.13 |
| 40 | 65 | 1.05 \pm 0.07 | 1.40 \pm 0.09 |
| 66 | 100 | 1.00 \pm 0.05 | 1.13 \pm 0.06 |
| 101 | 144 | 0.99 \pm 0.04 | 0.88 \pm 0.04 |
| 145 | 198 | 0.89 \pm 0.05 | 0.60 \pm 0.04 |
| 199 | 263 | 0.90 \pm 0.08 | 0.45 \pm 0.04 |
| 264 | 338 | 1.08 \pm 0.10 | 0.40 \pm 0.04 |
| 339 | 425 | 1.08 \pm 0.11 | 0.30 \pm 0.03 |
| 426 | 525 | 0.98 \pm 0.13 | 0.20 \pm 0.03 |
| 526 | 637 | 0.69 \pm 0.17 | 0.11 \pm 0.03 |
| 638 | 762 | 0.18 \pm 0.23 | 0.02 \pm 0.03 |
| 763 | 901 | 0.73 \pm 0.27 | 0.07 \pm 0.02 |
| 902 | 1054 | 1.11 \pm 0.40 | 0.08 \pm 0.03 |
| 1055 | 1221 | 0.27 \pm 0.55 | 0.01 \pm 0.03 |
| 1222 | 1404 | 1.06 \pm 0.80 | 0.05 \pm 0.03 |
| 1405 | 1602 | 0.91 \pm 1.18 | 0.03 \pm 0.04 |
| 1603 | 1816 | -2.03 \pm 1.75 | -0.06 \pm 0.05 |
| 1817 | 2048 | -0.54 \pm 2.16 | -0.01 \pm 0.05 |

Notes. The \hat{A}^6 values are dimensionless, with \hat{A}^6 = 1 for a measured spectrum equal to the lensing potential power spectrum of our fiducial cosmology (described in Sect. 2). The final column lists bandpower estimates, averaged within each bin.

In the following subsections, we discuss the information that can be gleaned from the lensing potential power spectrum alone, or in conjunction with the Planck 2015 CMB power spectrum likelihood. Some of the results in this section summarize the parts of Planck Collaboration XIII (2016) that are directly related to lensing.

3.5.1. Constraints from CMB lensing alone

The measurement of the lensing power spectrum is obtained from the 4-point function of the observed CMB anisotropies, and hence depends in general on both the lensing and CMB power spectra (see Appendix C). However the latter are now well measured by Planck, and there is only a weak dependence on cosmological parameters given the observed spectra (mainly via different cosmological models changing foreground parameters, and hence the amplitude of underlying CMB power spectra). In this section we fix the CMB power spectra to a \Lambda CDM Planck TT+lowP best fit (i.e. using TT on all scales and low-\ell polarization), and consider the conditional probability of different cosmological parameters given only the lensing reconstruction power spectrum.

To avoid marginalizing over very unrealistic values of poorly constrained parameters, we adopt several well-motivated priors when considering constraints from lensing alone.

- The optical depth to reionization is fixed to \tau = 0.07, because lensing deflections are independent of reionization (and scattering and subsequent lensing from sources at reionization is negligible).
- The baryon density is given a Gaussian 1\sigma prior \Omega_b h^2 = 0.0223 \pm 0.0009, as measured independently from big bang nucleosynthesis models combined with quasar absorption line observations (Pettini & Cooke 2012).
- The scalar spectral index is given a broad prior \ns = 0.96 \pm 0.02; results are only weakly sensitive to this choice, within plausible bounds.
- A top-hat prior is used for the reduced Hubble constant, 0.4 < h < 1. This limits the extent of the parameter degeneracy, but does not affect the results over the region of interest for joint constraints.

In addition to the priors above, we adopt the same sampling priors and methodology as Planck Collaboration XIII (2016)\textsuperscript{8}, using CosmoMC and camb for sampling and theoretical predictions (Lewis & Bridle 2002; Lewis et al. 2000). In the \Lambda CDM model, as well as \Omega_m h^2 and \ns, we sample \As, \Omega_b h^2, and the (approximate) acoustic-scale parameter \theta_M. Alternatively, we can think of our lensing-only results as constraining the sub-space of \Omega_m, \H, and \sigma_8. Figure 7 shows the corresponding constraints from CMB lensing, along with tighter constraints from combining with additional external baryon acoustic oscillation (BAO) data, compared to the constraints from the Planck

\textsuperscript{8} For example, we split the neutrino component into approximately two massless neutrinos and one with \sum m_\nu = 0.06 eV, by default.
CMC power spectra. The contours overlap in a region of acceptable Hubble constant values, and hence are compatible. To show the multi-dimensional overlap region more clearly, the red contours show the lensing constraint when restricted to a reduced-dimensionality space with $b_{MC}$ fixed to the value accurately measured by the CMB power spectra; the intersection of the red and black contours gives a clearer visual indication of the consistency region in the $\Omega_m-\sigma_8$ plane.

The lensing-only constraint defines a band in the $\Omega_m-\sigma_8$ plane, with the well-constrained direction corresponding approximately to the constraint

$$\sigma_8 \Omega_m^{0.25} = 0.591 \pm 0.021 \quad \text{(lensing only; 68%)}.$$  \hfill (13)

This parameter combination is measured with approximately 3.5% precision.

The dependence of the lensing potential power spectrum on the parameters of the $\Lambda$CDM model is discussed in detail in Appendix E; see also Pan et al. (2014). Here, we aim to use simple physical arguments to understand the parameter degeneracies of the lensing-only constraints. In the flat $\Lambda$CDM model, the bulk of the lensing signal comes from redshift ($z > 0.5$) where the Universe is mostly matter-dominated (so potentials are nearly constant), and from lenses that are still nearly linear. For fixed CMB (monopole) temperature, baryon density, and $n_s$, in the $\Lambda$CDM model the broad shape of the matter power spectrum is determined mostly by one parameter, $k_{eq} \equiv \ell_{eq} H_{eq} \propto \Omega_m h^2$.

The matter power spectrum also scales with the primordial amplitude $A_s$; keeping $A_s$ fixed, but increasing $k_{eq}$, means that the entire spectrum shifts sideways so that lenses of the same typical potential depth $\Psi_{\text{lens}}$ become smaller. Theoretical $\Lambda$CDM models that keep $\ell_{eq} \equiv k_{eq} \ell_{eq}$, fixed will therefore have the same number (proportional to $k_{eq}$) of lenses of each depth along the line of sight, and distant lenses of the same depth will also maintain the same angular correlation on the sky, so that the shape of the spectrum remains roughly constant. There is therefore a shape and amplitude degeneracy where $\ell_{eq} \propto k_{eq}$ constant, $A_s \approx$ constant, up to corrections from sub-dominant changes in the detailed lensing geometry, changes from late-time potential decay once dark energy becomes important, and nonlinear effects. In terms of standard $\Lambda$CDM parameters around the best-fit model, $\ell_{eq} \propto \Omega_m^{0.8} h^2$, with the power-law dependence on $\Omega_m$ only varying slowly with $\Omega_m$; the constraint $\ell_{eq} \propto \Omega_m^{0.8} h = \text{constant}$ defines the main dependence of $H_0$ on $\Omega_m$ seen in Fig. 7.

The argument above for the parameter dependence of the lensing power spectrum ignores the effect of baryon suppression on the small-scale amplitude of the matter power spectrum (e.g. Eisenstein & Hu 1998). As discussed in Appendix E, this introduces an explicit dependence of the lensing power spectrum on $\Omega_m h^2$. However, since the parameter dependence of $\ell_{eq}$ is close to $(\Omega_m h^2)^{1/2}$, we can still think of $A_s$ and $\ell_{eq}$ as giving the dominant dependence of $C_{l}^{\phi\phi}$ on parameters.

In practice the shape of $C_{l}^{\phi\phi}$ is not measured perfectly, and there is also some degeneracy between $A_s$, which directly increases the amplitude, and $\ell_{eq}$, which increases the amplitude via an increase in the number of lenses along the line of sight and indirectly by changing the depth of potential wells of a given size. On small angular scales with $L > \ell_{eq}$, $\Psi_{\text{lens}}$ becomes a strong function of scale due to the fall-off in the matter power spectrum for $k > k_{eq}$. For a fixed comoving scale, the amplitude increases if $k_{eq}$ becomes larger since this reduces the amount of decay of the gravitational potential during radiation domination (Pan et al., 2014), and also lessens the impact of baryon suppression. This increases the dependence of $C_{l}^{\phi\phi}$ on $\ell_{eq}$ beyond the linear scaling from the number of lenses, and on small scales $L^2 C_{l}^{\phi\phi} \propto A_s^{1+n_{eq}}$, where $n_L > 1$ determines the actual $L$-dependence (inherited from the scale dependence of the small-scale matter power spectrum and the effect of baryon suppression). For Planck, the lensing spectrum is best measured at $L \sim 200$: since $n_L$ increases with $L$, making $A_s$ larger and $\ell_{eq}^{1+n_{eq}}$ smaller to keep the power constant at $L \sim 200$ leads to an increase in power at lower $L$ and a decrease in power at higher $L$ (see Fig. 8). The lensing amplitude at the peak $L \approx 50$ (and hence the rms deflection angle), is therefore slightly anti-correlated with the amount of small-scale power (as shown by the lensing only constraint in Fig. 9; see discussion below).

The actual three-parameter constraint is close to

$$A_s \left(\Omega_m^{0.8} h^2\right)^{2.3} \propto A_s \ell_{eq}^{2.3} \propto \text{constant} \pm 3\%,$$  \hfill (14)
which, in terms of $A_\tau$, $\ell_{eq}$, and $h$, is independent of $h$, as expected, given the parameter dependence of $C_{\ell}^{\phi\phi}$ discussed above. Recall that the fairly high power-law dependence with $\ell_{eq}$ comes from the rapid fall-off in the matter power spectrum with scale (and from the dependence on $\Omega_m h^2$ of baryon suppression effects). Approximately the same scaling will also affect other measures of small-scale power, for example $\sigma_8$. From Eq. (E.20), for parameter values of interest $\sigma_8$ and $A_\tau$ are related approximately by

$$
\sigma_8^2 \Omega_m^{0.6} \propto A_\tau \ell_{eq}^{3.5}.
$$

(15)

The parameter direction $\sigma_8 \Omega_m^{3/5} = \text{constant}$ is therefore close to the direction in which $A_\tau$ (and hence $L^2 C_{\ell}^{\phi\phi}$) is constant for fixed $\ell_{eq}$. The combination $A_\tau \ell_{eq}^{3.5}$ is also not far from $A_\tau \ell_{eq}^{1.7}$ of Eq. (14), which remains nearly constant as $\ell_{eq}$ and $A_\tau$ vary within their degeneracy. The $\sigma_8 \Omega_m^{1/4}$ degeneracy direction shown in Fig. 7 and the constraint of Eq. (13) are therefore significantly tighter than the corresponding constraint in the $A_\tau-\Omega_m$ plane, because $\sigma_8$ absorbs some of the degeneracy between matter density and primordial power at a fixed scale.

In terms of $\sigma_8$, a well-determined three-parameter combination is

$$
\sigma_8 \Omega_m^{3/5} (\Omega_m h^2)^{0.37} = 1.228 \pm 0.028 \text{ (lensing only; 68%)},
$$

(16)

at 2.3\% precision. In terms of power, the 4.6\% precision of this square of this parameter combination is still weaker than the template amplitude measurement of Eq. (11a). This is expected, given the approximation of a power-law fit, and that in Eq. (16) we have also marginalized over variations in $n_s$ and $\Omega_b h^2$. Marginalizing out the $\Omega_b h^2$ dependence in Eq. (16), the constraint on $\sigma_8 \Omega_m^{0.25}$ of Eq. (13) is further degraded to 3.5\%.

We can also quantify the amplitude of lensing by the rms deflection angle $\langle d^2 \rangle^{1/2}$, which we define via

$$
\langle d^2 \rangle \equiv \sum_{L=2}^{2000} \frac{(2L + 1)}{4\pi} L(L + 1) C_{L}^{\phi\phi}.
$$

(17)

This is approximately an integral over all scales. It is proportional to the primordial fluctuation power $A_s$, and scales roughly as a power law in $\ell_{eq}$. A simple argument based on the number of lenses along the line of sight would suggest $\langle d^2 \rangle \propto A_s \ell_{eq}$. This is almost correct, in particular the strong dependence of $C_{\ell}^{\phi\phi}$ on $\ell_{eq}$ integrates out when forming $\langle d^2 \rangle$ (see Eq. (E.4)). However, there is a further dependence on $\Omega_m$ and $h$ due to the effects of baryon suppression and the late-time decay of the gravitational potential, which is important at low-$\ell$; we find approximately

$$
\langle d^2 \rangle \propto A_s \ell_{eq} (\Omega_m h^2)^{1/5},
$$

(18)

by taking derivatives of $C_{\ell}^{\phi\phi}$ computed with CAMB. The rms is measured at 2.4\% precision to be

$$
\langle d^2 \rangle^{1/2} = (2.46 \pm 0.06) \text{arcmin (lensing only; 68%)},
$$

(19)

So far we have only considered a CDM model. However, for most generalizations that are close to CDM, the same scaling also approximately holds, and the constraint direction of $\sigma_8 \Omega_m^{0.25}$ is only weakly model-dependent, although the centroid can shift and also the relation between $\Omega_m$ and $h$. For example, for a model with massless sterile neutrinos parameterized by $N_{\text{eff}}$, the constraint of Eq. (13) is virtually identical, but parameter samples shift towards systematically higher $\Omega_m h^2$ than in CDM to keep the equality scale roughly constant when $N_{\text{eff}} > 3.046$. As a further example, in models with three massive (active) neutrinos with degenerate masses, marginalizing over the mass of massive neutrinos shifts the constraint slightly down, giving

$$
\sigma_8 \Omega_m^{0.25} = 0.566 \pm 0.025 \text{ (ACDM+Simf, lensing only; 68%)},
$$

(20)

Note that in the neutrino case non-linear corrections from halofit are unreliable away from $\Omega_m \approx 0.3$, which can affect this result at around the 0.01 level.

3.5.2. Joint parameter constraints

The small-scale CMB power spectra are only directly sensitive to the primordial power $A_s$, after damping by reionization, through the combination $A_s e^{-\tau}$. However lensing also smooths the power spectra at the several percent level, which allows the amplitude $A_s$ to be constrained via the amplitude of the smoothing effect (and a small transfer of power to small scales), even when $\tau$ is not constrained precisely.\^10 The lensing smoothing can be thought of as approximately a convolution of the deflection angle power spectrum with the CMB power spectrum, and hence is mostly sensitive to scales around the peak of the deflection angle power spectrum. We can quantify the amplitude

\^10 The smoothing effect of lensing in the Planck TT power spectrum is detected at around $10\sigma$ (Planck Collaboration XIII 2016). It is also detected at lower significance in the polarization power spectra.
of the smoothing approximately by the mean-squared deflection angle \((d^2)\) from Eq. (17).

The CMB power spectra determine the angular acoustic scale \(\theta_a\) very accurately, which can be used to eliminate the \(h\)-dependence of \((d^2)\), considered as a function of \(\sigma_8\), \(h\), and \(\Omega_m\). Empirically, we find that in the \(\Lambda\)CDM model the CMB power spectrum posterior has \((d^2)^{1/2} \propto \sigma_8 \Omega_m^{0.25}\), close to the \(\sigma_8 \Omega_m^{0.25}\) combination measured more directly by lensing reconstruction\(^{1}\); see Fig. 9. For a flat \(\Lambda\)CDM model the constraint on the latter parameter is

\[
\sigma_8 \Omega_m^{0.25} = 0.621 \pm 0.013 \quad \text{(Planck TT+lowP; 68%)}. \tag{21}
\]

The amplitude information here is coming from a combination of the lensing smoothing effect, and information from the CMB power spectrum amplitude combined with the constraint on \(\tau\) from low-\(\ell\) polarization. The lensing-only constraint of Eq. (15) is slightly weaker, but somewhat lower than this (consistent at just over 1\(\sigma\)); see Fig. 7. The combined constraint is tighter and restricted to slightly lower values compared to the Planck power spectrum measurement:

\[
\sigma_8 \Omega_m^{0.25} = 0.609 \pm 0.007 \quad \text{(Planck TT+lowP+lensing; 68%)}. \tag{22}
\]

Here, for joint constraints we no longer apply the additional priors used in the lensing-only analysis, and allow for the full linearized CMB power spectrum dependence of the lensing renormalization in the likelihood. The overlap region in the \(\Omega_m-\sigma_8\) plane also excludes some of the lower Hubble constant values allowed by the power spectrum data, giving a slight shift of the mean in the direction of the higher values preferred by some local measurements (Riess et al. 2011; Humphreys et al. 2013; Efstathiou 2014), as shown in Fig. 9. This could be an additional indication that the lower Hubble constant values from the CMB spectra are at least partly a random statistical fluctuation, in which case joint constraints should be usefully closer to the truth. The joint constraint gives

\[
H_0 = (67.8 \pm 0.9) \text{ km s}^{-1} \text{ Mpc}^{-1} \quad \text{(Planck TT+lowP+lensing; 68%)}, \tag{23}
\]

about 0.5\(\sigma\) higher than Planck TT+lowP alone, and in good agreement with the constraint from Planck TT+lowP+BAO data.

The measurement of the fluctuation amplitude from lensing, combined with the amplitude of the CMB power spectrum on intermediate and small scales, can also be used to constrain the reionization optical depth independently of low-\(\ell\) polarization. The CMB power spectra themselves provide a weak constraint via the lensing smoothing, which is significantly improved by using the additional information in the lens reconstruction. Figure 10 shows the constraint when Planck CMB temperature power spectrum constraints, without low-\(\ell\) polarization, are combined with the lensing likelihood. We find

\[
\tau = 0.070 \pm 0.024 \quad \text{(Planck TT+lensing; 68%)}, \tag{24}
\]

\(^{1}\) The empirical parameter dependence of \((d^2)^{1/2}\) on \(\Omega_m\) for samples from the CMB power spectrum posterior differs from what one would infer from Eqs. (15) and (18). These equations assume fixed \(n_s\) and \(\Omega_b h^2\), which is appropriate for our lensing-only constraints with the assumed priors, but for the CMB samples the fractional dispersion in \(n_s\) and \(\Omega_m\) is comparable to that in \(\sigma_8\), \(h\), and \(\Omega_m\) and so cannot be neglected.
the joint constraint is consistent with a flat universe ($\Omega_K = 0$). The improvement in error bars is dramatic, with the CMB power spectra giving

$$\Omega_K = -0.052^{+0.03}_{-0.02} \quad (\text{Planck TT+lowP; 68%}),$$

and the joint constraint

$$\Omega_K = -0.005^{+0.009}_{-0.007} \quad (\text{Planck TT+lowP+lensing; 68%}).$$

This measurement of curvature from the CMB alone has sub-percent precision. External data can also break the geometric degeneracy; the joint BAO contours are shown in Fig. 11 and are consistent. Since the BAO results are independent of fluctuation amplitude, the lensing-reconstruction joint constraint adds significant additional information, shrinking the allowed parameter space compared to Planck+BAO alone.

In addition to curvature, CMB results for dark energy models also suffer a geometric degeneracy, and lensing results can similarly provide additional information; for applications to a variety of dark energy models and modified gravity see Planck Collaboration XIV (2016). Lensing can also improve the limits on many other parameters, though in a less dramatic fashion if the CMB power spectra are themselves able to place a good constraint. Full results for many combinations are available online in the grid tables and parameter chains12.

3.5.4. Sensitivity to \(L\)-range

The results in the previous subsections were obtained using our conservative baseline likelihood, with eight bins in the multipole range $40 \leq L \leq 400$. We now briefly discuss how joint constraints from the Planck power spectra and lensing change for different cuts of the lensing data.

Using the first nine bins of the aggressive binning described in Table 1, covering $8 \leq L \leq 425$, results from

\[ \sum m_\nu < 0.145 \text{ eV} \quad (\text{Planck TT+lowP+lensing+BAO; 68%}), \]

but using the full $L$-range the posterior peaks slightly (though not significantly) away from zero:

\[ \sum m_\nu = 0.16^{+0.08}_{-0.11} \text{ eV} \quad (\text{Planck TT+lowP+aggressive lensing +BAO; 68%}). \]

As discussed in Sect. 4, the multipole range $300 \leq L \leq 900$ may be unreliable, specifically due to the failure of the $TTTT$ curl test. A priority for future analysis should be to develop a trustworthy likelihood over the full multipole range.

4. Consistency and null tests

With multiple frequency bands and the combination of both temperature and polarization data, the Planck full-mission data set provides many opportunities to test internally the lens reconstruction. In this section we present consistency and null tests for the MV lensing potential presented in Sect. 3.

For consistency tests, we perform lens reconstruction using different data/analysis choices than those used for our fiducial MV reconstruction, obtaining an alternative measurement of the lensing potential power spectrum $C_L^{\phi\phi} |_{\text{test}}$, which we then compare to the baseline MV reconstruction. To compare the test and MV reconstructions more quantitatively, we use a $\chi^2$ statistic for the difference between the bandpowers, which is calculated as

\[ \chi^2_{\text{test}} = B_L^i \left( C_L^{\phi\phi} |_{\text{test}} - C_L^{\phi\phi} |_{\text{MV}} \right) \Sigma_{\text{test}}^{-1} \Sigma_{\text{test}} \left( C_L^{\phi\phi} |_{\text{test}} - C_L^{\phi\phi} |_{\text{MV}} \right). \]

where $\Sigma_{\text{test}}$ is a diagonal covariance matrix for the difference obtained from simulations of the two lens reconstructions. We compare this $\chi^2$ to a set of values determined from simulations, to estimate a PTE. We do this for several possible choices of binning: a single bin for either the conservative ($40 \leq L \leq 400$) likelihood range, or the low-$L$ and high-$L$ ranges (8 $\leq L < 40$ and 400 $< L \leq 2048$, respectively) that are additionally included in the aggressive likelihood; and the individual bins that span the conservative and high-$L$ multipole ranges. The single-bin tests check for broad-band differences between the two reconstructions, while the multiple-bin tests provide a more sensitive check on shape differences or inaccuracies in the error bars for the two reconstructions.

We group the consistency tests as follows, based on the issues which they are designed to probe.

- “Individual estimator crosses” that test for consistency among the five quadratic estimators used to measure $\phi$.
- “Foregrounds tests” that use single-frequency maps or different component-separation methods to test for possible foreground contamination.

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12 http://www.cosmos.esa.int/web/planck/pla
Planck Collaboration: Planck 2015 results. XV.

Fig. 12. Summary of lensing power spectrum consistency tests. Upper panel: comparison of the lensing power spectra for the baseline MV reconstruction and several alternative estimators with different data/analysis choices discussed in the text. Lower panels: detailed plots of the difference between each alternative reconstruction and the MV power spectrum, in units of the MV error bars ($\sigma$). The $\pm 1\sigma$ band is indicated in grey. Error bars on individual data points give the standard deviation of the expected difference from the MV reconstruction for each case, estimated using Monte Carlo simulations.

- “Noise tests” that check that our power spectrum estimates are not biased by instrumental noise, using cross-estimators designed to avoid noise biases.
- “Analysis choices” that test our sensitivity to choices such as the unmasked sky fraction and multipole range.
- “Bias-hardened estimators” in which we use lensing estimators that should be less sensitive to certain systematic and instrumental effects.

In Fig. 12 we plot several of the most important consistency tests, as well as their difference from the MV reconstruction. The $\chi^2$ values associated with the full set of consistency tests are presented in Table 2. The tests themselves are described in more detail in Sects. 4.1–4.5.

In addition to consistency tests, we also perform several null tests in which we apply our lens reconstruction and power spectrum estimation procedure to maps that have been differenced.
### Table 2. Amplitude fits for a suite of alternative reconstructions of the lensing power spectrum and comparison to the baseline MV lensing power spectrum estimate.

<table>
<thead>
<tr>
<th></th>
<th>Low-L</th>
<th>Consistency test summary values</th>
<th>High-L</th>
</tr>
</thead>
<tbody>
<tr>
<td>8 ≤ L &lt; 40</td>
<td>40 ≤ L ≤ 400</td>
<td></td>
<td>400 ≤ L ≤ 2048</td>
</tr>
<tr>
<td>( \hat{A} \pm \sigma_A )</td>
<td>( \Lambda_{\text{MV}} \pm \sigma_A )</td>
<td>( \Lambda_{\text{MV}} \pm \sigma_A )</td>
<td>( \Lambda_{\text{MV}} \pm \sigma_A )</td>
</tr>
<tr>
<td>( P_{\text{E}0\text{mv}} )</td>
<td>( P_{\text{E}0\text{mv}} )</td>
<td>( P_{\text{E}0\text{Bsm}} )</td>
<td>( P_{\text{E}0\text{mv}} )</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>1.04 ± 0.111</td>
<td>0.99 ± 0.032</td>
<td>0.72 ± 0.096</td>
</tr>
<tr>
<td>( \Delta T \times \Delta E )</td>
<td>1.00 ± 0.108</td>
<td>0.96 ± 0.039</td>
<td>0.62 ± 0.273</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>0.87 ± 0.149</td>
<td>1.02 ± 0.055</td>
<td>1.61 ± 0.448</td>
</tr>
<tr>
<td>( \Delta T \times \Delta E )</td>
<td>1.10 ± 0.308</td>
<td>0.98 ± 0.085</td>
<td>1.11 ± 0.270</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>0.89 ± 0.194</td>
<td>1.06 ± 0.079</td>
<td>3.60 ± 1.472</td>
</tr>
<tr>
<td>( \Delta E \times \Delta E )</td>
<td>1.07 ± 0.209</td>
<td>0.99 ± 0.085</td>
<td>0.44 ± 0.121</td>
</tr>
<tr>
<td>( \Delta E \times \Delta E )</td>
<td>0.64 ± 0.424</td>
<td>1.01 ± 0.157</td>
<td>0.78 ± 0.235</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>0.49 ± 0.631</td>
<td>0.89 ± 0.180</td>
<td>1.41 ± 0.575</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>1.13 ± 0.407</td>
<td>0.74 ± 0.189</td>
<td>4.50 ± 5.723</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>1.53 ± 0.664</td>
<td>0.64 ± 0.247</td>
<td>4.13 ± 2.332</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>2.24 ± 0.837</td>
<td>0.72 ± 0.297</td>
<td>5.24 ± 4.313</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>−0.98 ± 1.212</td>
<td>1.28 ± 0.463</td>
<td>3.15 ± 4.927</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>0.58 ± 0.189</td>
<td>1.40 ± 0.522</td>
<td>0.42 ± 0.174</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>0.87 ± 2.954</td>
<td>0.61 ± 0.397</td>
<td>0.34 ± 0.795</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>13.790 ± 8.353</td>
<td>−2.809 ± 2.189</td>
<td>−0.532 ± 8.268</td>
</tr>
<tr>
<td>( \Delta T \times \Delta T )</td>
<td>0.958 ± 0.097</td>
<td>0.98 ± 0.033</td>
<td>1.028 ± 0.193</td>
</tr>
<tr>
<td>Pol. only (^b)</td>
<td>1.252 ± 0.350</td>
<td>0.761 ± 0.145</td>
<td>1.787 ± 1.598</td>
</tr>
<tr>
<td>Foreground tests</td>
<td>94.9 ± 0.096</td>
<td>1.024 ± 0.034</td>
<td>0.583 ± 0.122</td>
</tr>
<tr>
<td>217 GHz</td>
<td>0.941 ± 0.092</td>
<td>0.963 ± 0.034</td>
<td>0.929 ± 0.109</td>
</tr>
<tr>
<td>NILC</td>
<td>0.977 ± 0.073</td>
<td>0.988 ± 0.026</td>
<td>0.877 ± 0.082</td>
</tr>
<tr>
<td>SEVEM TT (^a)</td>
<td>1.015 ± 0.010</td>
<td>0.800 ± 0.011</td>
<td>0.758 ± 0.094</td>
</tr>
<tr>
<td>Commander TT (^a)</td>
<td>0.899 ± 0.124</td>
<td>0.901 ± 0.013</td>
<td>0.924 ± 0.011</td>
</tr>
<tr>
<td>( \delta A_{\text{sys}} = 0.8 )</td>
<td>1.009 ± 0.004</td>
<td>0.987 ± 0.025</td>
<td>0.761 ± 0.086</td>
</tr>
<tr>
<td>( \delta A_{\text{sys}} = 0.6 )</td>
<td>1.000 ± 0.007</td>
<td>0.991 ± 0.026</td>
<td>0.777 ± 0.091</td>
</tr>
<tr>
<td>( \text{SMICA} \times \text{353 GHz} )</td>
<td>1.000 ± 0.013</td>
<td>0.970 ± 0.048</td>
<td>0.946 ± 0.134</td>
</tr>
<tr>
<td>( \text{SMICA} \times \text{353 GHz pol. only} )</td>
<td>0.277 ± 0.096</td>
<td>0.264 ± 0.451</td>
<td>7.04 ± 5.895</td>
</tr>
<tr>
<td>Noise tests</td>
<td>0.949 ± 0.077</td>
<td>0.999 ± 0.027</td>
<td>0.755 ± 0.095</td>
</tr>
<tr>
<td>432 ± 217 GHz</td>
<td>1.028 ± 0.081</td>
<td>0.981 ± 0.029</td>
<td>0.827 ± 0.090</td>
</tr>
<tr>
<td>( \text{HMI} \times \text{HM2} )</td>
<td>0.989 ± 0.070</td>
<td>0.994 ± 0.026</td>
<td>0.779 ± 0.084</td>
</tr>
<tr>
<td>Noise scaling (^{11})</td>
<td>1.136 ± 0.104</td>
<td>0.990 ± 0.030</td>
<td>0.718 ± 0.153</td>
</tr>
<tr>
<td>Apodization</td>
<td>1.009 ± 0.088</td>
<td>0.985 ± 0.026</td>
<td>0.753 ± 0.097</td>
</tr>
<tr>
<td>Bias-hardened estimators</td>
<td>( \hat{A}^{TT} )</td>
<td>1.117 ± 0.120</td>
<td>0.968 ± 0.039</td>
</tr>
<tr>
<td>( \hat{A}^{TT} )</td>
<td>1.068 ± 0.116</td>
<td>0.983 ± 0.036</td>
<td>0.651 ± 0.103</td>
</tr>
<tr>
<td>( \hat{A}^{TT} )</td>
<td>1.082 ± 0.118</td>
<td>0.983 ± 0.038</td>
<td>0.617 ± 0.107</td>
</tr>
</tbody>
</table>

**Notes.** The measured overall lensing amplitude \( \hat{A} \) is given for the conservative (40 ≤ L ≤ 400) multipole range, as well as the lower signal-to-noise bins on either side. The difference between the test amplitudes and the corresponding MV result (with amplitudes of \( \Lambda_{\text{MV}} = 1.023 ± 0.085 \), \( \Lambda_{\text{MV}} = 0.978 ± 0.025 \), and \( \Lambda_{\text{MV}} = 3.848 ± 0.078 \) for the three multipole ranges) is given as \( \Delta \Lambda_{\text{MV}} \). The quoted uncertainty on the difference is obtained using Monte Carlo simulations and therefore accounts for any correlations between the MV and test estimates. The PTE values are estimated by comparison of the \( \chi^2 \) values to the distribution from simulations. PTEs are given for the square of the overall amplitude difference, as well as for the \( \chi^2 \) of the difference between the handpowers of the binned spectra. \(^a\) Comparison is against the SMICA polarization-only reconstruction, rather than the MV reconstruction. \(^b\) Comparison is against the SMICA TT reconstruction, rather than the MV reconstruction. \(^{11}\) This test only involves a change in the simulations used to perform the debiasing corrections and estimate error bars. We have quoted the scatter between the two sets of simulations, although this is an upper limit on the difference that we should actually see for the data.
Table 3. Null tests for a set of lens reconstructions based on maps that should contain no CMB signal, or estimators that should be orthogonal to the lensing signal.

<table>
<thead>
<tr>
<th>Null test summary values</th>
<th>Low-L 8 ≤ L &lt; 40</th>
<th>Conservative 40 ≤ L ≤ 400</th>
<th>High-L 400 &lt; L &lt; 2048</th>
</tr>
</thead>
<tbody>
<tr>
<td>Half-ring difference</td>
<td>(\hat{\chi}_{\text{null}}^2)</td>
<td>(\hat{\chi}_{\text{null}}^2)</td>
<td>(\hat{\chi}_{\text{null}}^2)</td>
</tr>
<tr>
<td>(\hat{\psi}<em>{\text{TE}}\times\hat{\psi}</em>{\text{TE}}) (×1)</td>
<td>0.78</td>
<td>0.03</td>
<td>0.54</td>
</tr>
<tr>
<td>(\hat{\psi}<em>{\text{EB}}\times\hat{\psi}</em>{\text{EB}}) (×0.5)</td>
<td>0.63</td>
<td>0.39</td>
<td>0.18</td>
</tr>
<tr>
<td>(\hat{\psi}<em>{\text{TT}}\times\hat{\psi}</em>{\text{TT}}) (×10)</td>
<td>0.87</td>
<td>0.59</td>
<td>0.93</td>
</tr>
<tr>
<td>(\hat{\psi}<em>{\text{EE}}\times\hat{\psi}</em>{\text{EE}}) (×1)</td>
<td>0.13</td>
<td>0.66</td>
<td>0.41</td>
</tr>
<tr>
<td>(\hat{\psi}<em>{\text{MB}}\times\hat{\psi}</em>{\text{MB}}) (×15)</td>
<td>0.60</td>
<td>0.18</td>
<td>0.03</td>
</tr>
</tbody>
</table>

Notes. The measured overall lensing amplitude \(\hat{\chi}\) is given for the conservative (40 ≤ L ≤ 400) multipole range, as well as lower and higher multipoles. All should be consistent with zero. The PTE values are estimated by comparison of the \(\hat{\chi}\) values to the distributions from simulations. PTEs are given for the different of the overall amplitude from zero, as well as for the \(\chi^2\) of the difference from zero of the binned spectra.

Fig. 13. Power spectrum null tests discussed in Sect. 4.6. Curl null tests are denoted by \(\psi\). The individual null test power spectra have been scaled by the factors indicated in each panel for visual purposes. The fiducial (gradient-mode) lensing power spectrum is shown as the dashed line in each panel.

to remove sky signal, or use curl-mode lensing estimators that should have zero contribution from the gradient lensing potential. We check that the overall amplitude (for a template given by the fiducial theory spectrum \(C_{\ell}^{\text{fid}}\)) is consistent with zero, and also evaluate a \(\chi^2\) computed as

\[
\hat{\chi}_{\text{null}}^2 = \Sigma^{-1}_{\text{null}} \left[ \hat{\chi}_{\text{null}} \right]^T \Sigma^{-1}_{\text{null}} \left[ \hat{\chi}_{\text{null}} \right],
\]

where \(\Sigma_{\text{null}}\) is a diagonal covariance matrix measured from simulations of the particular null test. A PTE for each null test is computed by comparison to simulations in the same way as for the consistency tests. The null power spectra are plotted in Fig. 13, and summary statistics are given in Table 3. The null tests are discussed in more detail in Sect. 4.6.

4.1. Individual estimator crosses

With the foreground-cleaned temperature and polarization maps from SMCa, we can form five distinct lensing estimators, and 15 distinct auto- and cross-spectra. We plot all of these spectra individually in Fig. 14. The lensing amplitudes, as well as PTEs for the differences from the MV reconstruction are given in Table 2. There is a wide variation in the sensitivity of the different estimators, although we see high (greater than 5σ significance) detections of lensing in nine of the spectra.13

In the 2015 Planck data release, there is some evidence in the polarization power spectra for uncorrected systematic effects associated with temperature-polarization leakage (Planck Collaboration XIII 2016; Planck Collaboration XI 2016). Although we find no evidence for inconsistencies in the polarization-based lens reconstructions, the statistical errors are large and may obscure small systematic effects that are subdominant for the current analysis. Leakage from the CMB temperature anisotropies could impact the lensing analysis in several ways. First, the disconnected (Gaussian) part of the CMB 4-point function is modified, and this would affect both our mean-field and N0 corrections (see Appendices A.2 and A.3). These corrections are made using a mix of data and simulations, such that even if the simulations do not fully capture the leakage (note they do include leakage due to main-beam mismatch, for example), the additive bias to the lensing power spectrum is only second-order in any unmodelled effects. The second way that leakage can affect the lensing power spectrum is through changes to the lensing response functions in Eq. (A.13), which propagate through to normalization errors. This should be partially captured by a Monte-Carlo correction that we make to the lensing power spectrum. We have not attempted to quantify more carefully any residual effects of temperature-polarization leakage, since, ultimately, the polarization data carries relatively little weight in our final MV reconstruction.
As a further test, we combine several of the spectra with the highest S/N that only use estimator crosses (TT × EE, TT × TE, TT × EB, and EE × TE) into a single “crosses” estimator. This combination of cross-spectra has an order-of-magnitude smaller disconnected noise bias correction than the auto-spectrum of the fiducial MV estimator, with $\Delta C_8^{\text{MV}}$ _Crosses_ = 0.4, compared to $\Delta C_8^{\text{MV}}$ = 4.3 for the baseline MV lensing reconstruction. The results of this test are presented in Fig. 12 and Table 2. With the crosses power spectrum estimates we are able to test that the lensing amplitude is not biased by errors in the disconnected noise bias correction $\Delta C_8^{\text{MV}}$ at around the 2% level.

We also form a combined power spectrum estimate using only polarization data (the EE and EB estimators). This “Pol. only” estimator gives an amplitude estimate for the conservative multipole range of $A = 0.76 \pm 0.15$, a 5σ detection of lensing at a level consistent with the fiducial MV reconstruction.

### 4.2. Foreground tests

The SMICA CMB map that is used for our baseline MV reconstruction combines multiple frequency bands as a function of angular scale to reject foregrounds and to produce a single high-fidelity estimate of the CMB. We can also compare to a simpler approach based on individual frequency maps, which we clean by subtracting a dust template. Our baseline results in Planck Collaboration XVII (2014) were based on this approach using 143 GHz and 217 GHz data, with the 857 GHz map projected out during the filtering process as a dust template. In this work we repeat this analysis. We use 857 GHz as a dust template in temperature, and no cleaning in polarization. The results of this test are presented in Fig. 12 and Table 2 with the labels “143 GHz” and “217 GHz”, where it can be seen that we obtain lensing bandpowers that are consistent with the MV reconstruction for the conservative multipole range. In the simulations used to characterize these reconstructions, we add coloured Gaussian noise to the 143 GHz and 217 GHz temperature map simulations, designed to mimic the power of unsubtracted extragalactic foregrounds. For this we use the same source model used in Planck Collaboration XVII (2014). We also add Gaussian power to the polarization maps, designed to mimic the power due to Galactic dust contamination that can be seen on large angular scales. The resulting temperature and polarization simulations have power spectra that agree with those of the data at the percent level in power.

We can also test the stability of the lensing reconstruction to the choice of the SMICA component-separation method over other alternatives such as NILC, SEVEM, and Commander. The results of these comparisons are also shown in Table 2; all are very consistent with the SMICA-based reconstruction over the conservative multipole range, although the PTEs for the overall amplitude at high-$L$ (400 < $L$ ≤ 2048) are low for NILC and Commander. For the SEVEM and Commander analyses, polarization maps were not available at the high resolution used here ($N_{\text{side}} = 2048$), and so we have limited the comparisons to temperature only.

As a test of Galactic contamination, we also repeat our baseline analysis (performed on a fraction $f_{\text{sky}} = 0.7$ of the sky) using a more aggressive ($f_{\text{sky}} = 0.8$) and more conservative ($f_{\text{sky}} = 0.6$) Galactic mask. These masks are constructed by thresholding maps from the Planck high-frequency channels to achieve the desired sky fraction. In Table 2 we see that for both
alternative sky fractions we estimate lensing bandpowers that are consistent with the baseline.

As a final and very powerful test of high-frequency foreground contamination, we also perform a lens reconstruction in which we use the 353 GHz map as one of the four legs in the trispectrum used to estimate $C_{\ell}^{TT}$. Based on the approximate frequency scaling of dust (Gispert et al. 2000) we expect that any dust foreground bias $\sim \theta_{\text{dust}}$ to the lensing power spectrum estimate that we see in this test will scale roughly as

$$b_{\text{dust}}^{353} \approx \theta_{\text{dust}}^{17} \approx 20 b_{\text{dust}}^{143},$$

and so this should provide a sensitive test for such contamination in the SMICA map. The 353 GHz map has significantly more power than CMB+noise-only simulations, due primarily to Galactic dust, and temperature fluctuations from the CIB. To obtain reasonable error bars from the 353 GHz simulations used to characterize this reconstruction, as with the other reconstructions we add an isotropic Gaussian sky signal designed to match the simulation power spectrum to a smoothed power spectrum of the data. When filtering the 353 GHz map, we use the same parameters as used for the SMICA map (except for the beam transfer function), to have the same weight given to different angular scales in the quadratic lensing estimators. As with the 143 GHz and 217 GHz analyses, we project out the 857 GHz map as a dust template from the 353 GHz map in temperature, but make no attempt to subtract foreground contamination in polarization. As can be seen in Fig. 12 and Table 2, we obtain good agreement between the “one leg 353 GHz” reconstruction and the MV lensing power spectrum. Lensing power is detected at 20σ even using 353 GHz as one of the four estimator legs. The stability of the lens reconstruction in the presence of the foreground power at 353 GHz is very striking. As a further test for the stability of the polarization-based lens reconstructions, we can also use the “Pol. only” estimator from Sect. 4.1 with one leg given by the 353 GHz data. This reconstruction is very noisy, however we do obtain an overall amplitude as well as bandpowers that are consistent with the MV “Pol. only” reconstruction, as shown in Table 2.

4.3. Noise tests

As mentioned in Sect. 2, the FFP8 simulations on which our results are based underestimate the noise power by several percent at high $\ell$. We have dealt with this problem in our simulations by including isotropic Gaussian power to make up for the difference. This neglects the fact that the missing noise component most likely has an anisotropic distribution (which could lead to errors in the mean-field subtraction), and could be non-Gaussian as well. To test that our treatment of the missing noise component is adequate, we have performed several tests. Most stringently, we form quadratic lensing estimates using cross-correlation between pairs of maps with independent noise realizations. We construct both a frequency cross 143 GHz $\times$ 217 GHz, as well as a cross between the first and second halves of the mission data. These estimators have no noise mean field, and the lensing power spectrum estimated from them is insensitive to non-Gaussianity of the instrument noise. The results of this analysis are given in Table 2, where we see results that are consistent with the MV power spectrum estimate.

As an additional test, we have recreated the analysis of Planck Collaboration XVII (2014), where the missing noise power was accounted for by scaling the noise component of the FFP8 simulations by a small factor. This approach reasonably supposes that the anisotropy of the missing noise component is the same as that of the modeled component. This test is labelled “Noise scaling” in Table 2. Again, we see a power spectrum measurement with an amplitude that is very similar to the MV estimate (although we note that because the only difference in this test is the construction of the simulations that are used, the expected scatter between the “Noise scaling” and baseline approaches is difficult to determine).

4.4. Analysis choices

We test the stability of our reconstruction to several analysis choices, re-running the baseline analysis implementing the following changes one at a time.

- Use a high-pass filter in temperature and polarization with $\ell_{\text{min}} = 1000$ rather than $\ell_{\text{min}} = 100$, as adopted in the baseline analysis.
- Use a low-pass filter with $\ell_{\text{max}} = 1500$ rather than $\ell_{\text{max}} = 2048$, as used in the baseline analysis.
- Apodize the lensing convergence estimate before power spectrum estimation, following the analysis in Planck Collaboration XVII (2014). This change reduces the size of the Monte Carlo error, but does not produce a significant change in the lensing bandpower estimates.

The results of these tests are given in Table 2.

4.5. Bias-hardened estimators

Throughout this paper, we have used quadratic estimators optimized to detect the anisotropic covariance between CMB modes that is induced by (fixed) lensing. These estimators are biased by non-lensing sources of such mode coupling, including effects such as galaxy and point-source masking, inhomogeneous instrumental noise, and beam asymmetries. As discussed in Sect. 2, we estimate these biases using Monte Carlo simulations and subtract them as a “mean-field” contribution from the estimated lensing potential (as in Eq. (A.17)). A complementary method, which is less reliant on the quality of the Monte Carlo simulations used to perform this debiasing, is the bias-hardening procedure advocated by Namikawa et al. (2013). In this approach, one constructs a new quadratic lensing estimator that is orthogonalized to worrisome sources of bias. The construction of these estimators is reviewed in Appendix B. In the analysis here we focus on temperature-only bias-hardened estimators to investigate the effects of bias-hardening, because mean-field corrections are much smaller in polarization than in temperature. We denote these estimators as $\mathcal{F}_{TT} - X$, where $X$ denotes the bias(es) that are orthogonalized against. Here, we consider the biases due to masking ($M$), inhomogeneous noise ($N$), and the shot noise due to unresolved point sources ($S$). As discussed in Osborne et al. (2014), the source-hardened estimator $\mathcal{F}_{TT} - S$ also reduces contamination from effects such as the cross-correlation between the unresolved sources and the lensing potential.

Mean-field corrections are largest at low-$L$ (the power spectrum of the mean-field correction is larger than the fiducial lensing potential power spectrum for $L \lesssim 100$), and so we begin by looking closely at these scales. In Fig. 15 we plot the power spectra of the mean-fields for each of the bias-hardened estimators. In the lower panel we also plot the difference between each estimator and the non-bias-hardened $TT$ estimator. For $L \gtrsim 8$, we see agreement between all of the bias-hardened estimators and the non-bias-hardened reconstruction, regardless
of the much smaller mean-field corrections for the former. We have also used the bias-hardened estimators to measure the lensing power spectrum on smaller scales. In Fig. 12 and Table 2 we report lensing bandpower estimates using the bias-hardened estimators, which we find to be consistent with the baseline MV reconstruction for $L \leq 400$.

4.6. Null tests

In addition to the consistency tests described above, we have also performed several direct null tests, either using estimators or maps for which the lensing signal is expected to be zero. These are plotted in Fig. 13, and significance statistics for each are given in Table 3. We discuss them in more detail below.

Half-ring difference: the Planck observation strategy consists of scanning the sky in rings with an opening angle of approximately 85°. Each scan ring is observed for between 39 and 65 minutes, and then the satellite is repointed. This leads to a natural “half-ring” null test in which we run our standard lensing estimator on the half-difference of maps constructed from the first and second half of each scanning ring, which should consist entirely of noise. The PTE for the spectrum of the lensing reconstruction estimated from these half-ring difference maps is rather low for the conservative multipole range (see Table 3). Since the maps should be only noise by construction, to obtain an acceptable $\chi^2$ for the power spectra of reconstructions based on such maps requires very good modeling of the instrument noise. As can be seen from the factor of 200 scaling applied in Fig. 13, the impact on the reconstructed $C_L^{\phi\phi}$ of minor errors in the noise modeling that likely drive the low PTEs for the null test should be negligible.

Half-mission difference: noise that is correlated between rings (for example, due to glitches and destriping) can be nulled in the half-ring difference. Another differencing approach, which avoids this problem, is to divide the entire mission into two halves that are then differenced. This half-mission difference requires a slightly larger analysis mask, to account for the fact that the sky coverage is imperfect in both the first and second halves of the mission and some pixels are missed. The PTEs for the half-mission test are reasonable.

Curl-mode estimators: in general, deflections due to lensing can be decomposed into gradient ($\phi$) and curl ($\psi$) modes. The curl modes are expected to be consistent with zero for lensing by scalar perturbations at Planck reconstruction noise levels. We construct curl-mode power spectrum estimators, analogous to those used for the gradient lensing modes, and measure their power spectra. Full-sky weight functions for these estimators are presented in Namikawa et al. (2012). Note that although they do not couple directly to the lensing potential, the curl power spectra do receive “N1”-type contributions from the lensing trispectrum, which we estimate analytically and subtract as with the $C_L^{\phi\phi}$ power spectrum estimates. We do not make “PS” or “MC” corrections for the curl-mode power spectra. In Table 3 we report estimated amplitudes for a signal with the shape of the fiducial lensing power spectrum $C_L^{\phi\phi, \text{fid}}$ in the curl estimates. These amplitudes are all consistent with zero at better than $2\sigma$, with the exception of the $TT$ reconstruction over the high-$L$ range ($400 < L \leq 2048$). Indeed, there is mild evidence for a broad feature in the TT curl-mode reconstruction centred around $L \approx 500$ (see Fig. 13), which drives the low PTE of the TT curl-mode power spectrum over the high-$L$ multipole range. Considering $TT$ alone, the average curl power is low by about $2.9\sigma$ in the high-$L$ range, which is suggestive but not clearly unacceptable given the number of tests performed. Combined with the polarization reconstructions, which on their own appear consistent, the MV estimate remains low but only at $2\sigma$ over the high-$L$ range. Since the origin of this feature in the curl-mode spectrum is currently not understood, we have chosen to adopt the conservative multipole range for the baseline Planck 2015 lensing likelihood. This multipole range has little overlap with that covered by the curl-mode feature and was determined in advance of the curl result as it retains the majority of the S/N in the reconstruction.

5. Conclusions

We have presented a first analysis of gravitational lensing in the Planck 2015 full-mission data set. Using temperature and polarization data, we make the most powerful measurement of...
CMB lensing to date, with an overall significance of around 40σ. This paper is accompanied by the public release of the lensing potential map (as well as simulations that can be used to characterize the map, e.g. for cross-correlation analyses with external data) and a conservative lensing bandpower likelihood.

We have performed a number of tests to verify the internal consistency of the lens reconstruction, in particular we have tested the compatibility between individual temperature and polarization estimators, sensitivity to various analysis choices, and possible foreground contamination. Given Planck’s unique full-sky coverage, the full-mission maps are unlikely to be displaced as the most powerful data set for cosmological lensing analysis for the next several years (until high signal-to-noise E- and B-mode polarization measurements can be made on sufficiently large sky fractions). There is some room, however, for improvement on the analysis presented here. The inverse-variance filtering we have used does not account for variations in the instrument noise level across the sky, which could be incorporated to reduce the reconstruction noise (particularly in polarization, where instrumental noise dominates the signal). We have high-pass filtered all results to L ≥ 8, due to concerns about the lensing mean-field correction. A more aggressive analysis could potentially remove this cut. The conservative L ≤ 400 cut that we make in our baseline likelihood, imposed because of mild evidence of a broad feature in curl-mode tests beyond this range, might also be relaxed in a future analysis, although only with minor improvements in the statistical power of the likelihood.

The microwave background is a unique source for lensing studies, as the most distant and well-understood source plane that can be observed. The measurement presented here represents an integrated measurement of the total matter distribution in the entire observable Universe, and another powerful test of the veracity of the concordance ΛCDM cosmological model.

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References


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Appendix A: Pipeline

In this appendix we provide a more detailed technical overview of the pipeline that we use to extract lensing information from the Planck maps. Starting from input CMB sky maps the pipeline has the following three distinct analysis stages.

(1) Filtering. In this step, we perform a linear operation on the maps, which performs three tasks.
   a. Mask out regions of the sky that may be contaminated by point sources or residual foreground emission (primarily from the Galaxy).
   b. Deconvolve the instrumental beam and pixel transfer functions.
   c. Downweight noisy modes to allow later steps of the pipeline to produce optimal (in the minimum-variance sense) estimates of the lensing potential and its power spectrum.

Implementation details of the filtering procedure are described in Appendix A.1. The filtering step provides us with filtered CMB temperature and polarization multipoles denoted by $T_{\ell m}, E_{\ell m}, \text{and } B_{\ell m}$.

(2) Quadratic Estimators. The filtered CMB multipoles are then fed into quadratic estimators designed to extract the statistical anisotropy that is induced by lensing. These estimators are formed by summing over pairs of CMB fluctuations with a weight function optimized to detect the off-diagonal contributions to their covariance matrix from (fixed) lenses. This step produces maps of the estimated lensing potential $\hat{\phi}$, and is described in detail in Appendix A.2.

(3) Power Spectrum Estimation. To estimate the power spectrum of the lensing potential, we exploit the non-Gaussianity induced by lensing. At first order in the power spectrum $C_{\ell}^{TT}$, lensing generates a contribution to the connected 4-point function (or trispectrum) of the observed CMB, which we probe using auto- and cross-spectra of the 2-point quadratic estimators $\hat{\phi}$ obtained in step (2) of the pipeline. There are two complications here that are worth noting: the auto- and cross-spectra of the lensing estimators have noise biases even in the absence of lensing, which must be carefully estimated and subtracted; and non-Gaussian contamination from unresolved point sources can mimic lensing and bias the reconstructed power spectrum. Our approaches to these issues are described in Appendix A.3.

We now proceed to describe the individual steps of the lensing pipeline in more detail.

A.1. Filtering

At this stage in the pipeline, we have noisy maps of the CMB for both data and simulations. To construct minimum-variance estimates of the lensing potential and its power spectrum, we need to filter these maps to downweight noise-dominated modes, as well as to deconvolve the transfer functions due to the finite beam size and pixelization.

To derive the optimal (minimum-variance) filter to apply, we use a simplified model of the data given by

$$d_i = \sum_k Y_{i,k} s_k + n_i,$$

where $d_i$ is a single long vector containing the pixelized $T$, $Q$, and $U$ sky maps. The "pointing matrix" $Y_{i,k}$ takes an input sky signal (consisting of $T$, $E$, and $B$ multipoles collectively indexed by $k$), convolves it with a beam and pixel transfer function, and then performs a harmonic transform to map space. The harmonic modes of the sky itself are denoted as $s_k$, and $n_i$ is the map noise realization. We have used indices $i$ and $k$ for quantities in map-space and harmonic-space respectively. Explicitly, the indices $k$ and $i$ are mapped as

$$k \mapsto (X \in \{T, E, B\}, \ell, m) \quad \text{(A.2)}$$

$$i \mapsto (M \in \{T, Q, U\}, p), \quad \text{(A.3)}$$

where $p$ denotes a pixel. In this notation, we model the pointing matrix as

$$Y_{i,M}(X,m) = H_I B_I R_{X,m}^{T}(\hat{\eta}_p), \quad \text{(A.4)}$$

where $H_I$ is the HEALPix (Górski et al. 2005) $N_{\text{side}} = 2048$ pixel transfer function, $B_I$ is the effective beam transfer function for the map under consideration, and $R_{X,m}^{T}(\hat{\eta}_p)$ represents the real-valued spherical harmonic for field $X$ evaluated at the center of pixel $p$ in map $M$.

If both the sky signal and noise are Gaussian random fields (with covariance matrices $\langle s_k s_{k'}^* \rangle = S_k$ and $\langle n_i n_j^* \rangle = N_{ij}$, respectively) then the inverse-variance filtered estimates of the beam-deconvolved sky signal $\hat{s}$ are given in matrix notation by (Bunn et al. 1994)

$$\hat{s} = S^{-1} \left[ Y^T N^{-1} Y \right]^{-1} Y^T N^{-1} d.$$  

We use this equation to obtain the inverse-variance filtered $T$, $E$, and $B$ multipoles that are fed into our quadratic estimators, with one additional rescaling discussed below. The large bracketed matrix inverse is performed using conjugate descent with a multi-grid preconditioner, following Smith et al. (2007).

For the purposes of filtering, we use a simplified model for the covariance matrices of $s_k$ and $n_i$ in which they are diagonal in harmonic space and pixel space, respectively, with

$$S = \langle s_{(X,m)} s_{(X',m')}^* \rangle = \delta_{XX'} \delta_{\ell\ell'} \delta_{mm'} C_{\ell}^{XX'},$$

$$N = \langle n_{(M,p)} n_{(M',p')}^* \rangle = \delta_{MM'} \delta_{pp'} N_p^M.$$  

Here, $C_{\ell}^{XX}$ is the fiducial theoretical CMB power spectra and $N_p^M$ gives the noise variance for pixel $p$ in map $M$. With this diagonal assumption, we have ignored the cross-correlation $C_{\ell}^{TE}$ for the purposes of filtering. This is slightly sub-optimal, but has the advantage of allowing the temperature and polarization maps to be filtered independently. It also has the benefit of making analytical calculations for the estimator normalization simpler (see Sect. A.2). Also for simplicity, we take $N_p^M$ to be constant over all unmasked pixels and given by

$$N_p^M = \left( \frac{\pi N_{\text{pix}}^{\text{DB}}} {10,800} \right)^2 N_{\text{pix}} \frac{1} {4\pi} M_p.$$  

Here $N_{\text{pix}}^{\text{DB}}$ is the map noise level in $\mu K$ arcmin, $N_{\text{pix}}$ is the number of pixels in an $N_{\text{side}} = 2048$ HEALPix map, and $M_p$ is a mask map (which is zero for masked pixels, and unity otherwise). We always take the noise level to be the same for paired $Q/U$ maps. The mask term in Eq. (A.7) takes the noise level to infinity for masked pixels. In the filtering operation, which only involves the inverse of the noise matrix, this sets masked pixels to zero.
After evaluating Eq. (A.5), we have a set of inverse-variance filtered sky multipoles \( \bar{s}_{X,\ell,m} \) for \( X \in \{T, E, B\} \). The simplified CMB+white-noise model above results in a slightly sub-optimal filtering procedure since it ignores the following:

1. variations in the noise level across the sky due to the uneven hit distribution (the Elliptic poles receive more coverage than the equator for the Planck scan strategy);
2. scale-dependence of the noise power; and
3. foreground power.

In Planck Collaboration XVII (2014) we showed that the sub-optimality due to (1) for temperature-based lensing reconstruction was small (accounting for it properly was estimated to improve the uncertainty on the overall lensing amplitude by 4%). The loss will be slightly smaller here for temperature, owing to a more even hit distribution in the full-mission data than in the nominal-mission data used in Planck Collaboration XVII (2014). The loss will be larger for polarization, where noise is a more important part of the total error budget, although we have not evaluated the size of the potential degradation. To compensate for (2) and (3) we make a post-correction by rescaling the modes by a quality factor \( Q(X) \) as a function of scale to obtain final filtered multipoles that are closer to optimal. We denote these as \( \hat{X}_{\ell,m} \), and they are given by

\[
\hat{X}_{\ell,m} = Q^X \bar{s}_{X,\ell,m}.
\]  

These inverse-variance-filtered multipoles are the input to the lensing estimators described in Sect. A.2. The choice of \( Q^X \) is discussed below.

For analytical calculations, it is useful to have an approximation to the filtering procedure that is diagonal in harmonic space (i.e. neglecting the small amount of mode-mixing induced by the masking). In this approximation, the filtered multipoles are given by

\[
\hat{X}_{\ell,m} \approx F^X X_{\ell,m} + \hat{n}_{\ell,m},
\]

where \( \hat{n} \) is a noise realization and

\[
F^X = \frac{Q^X}{C^XX + N^XX}.
\]

Here, \( C^XX \) are the fiducial CMB auto-spectra used in Eq. (A.6), and the quantity

\[
N^XX = (H_0 B_s)^2 \left( \frac{n N_m}{10^{80}} \right)^2
\]

is the pixel- and beam-deconvolved noise power spectrum for the field \( X \), where \( N_m \) is the appropriate map white-noise level in \( \mu K \) arcmin. We choose \( Q^X \) such that

\[
\frac{f_{\text{sky}}}{2(L + 1)} \sum_m |\hat{s}_{X,\ell,m}|^2 \approx F^X,
\]

where \( f_{\text{sky}} = \sum_p M_p / N_{\text{pix}} \) is the unmasked sky fraction and the approximation sign indicates a smoothing over multipoles.

### A.2. Quadratic estimators

Ensemble averaging over a fixed realization of the CMB lensing potential, the CMB covariance matrix acquires off-diagonal elements given by

\[
\Delta(X_{\ell,m}, Z_{m',\ell'}) = \sum_{LM} (-1)^{M} \left( \frac{l_1}{m_1} \frac{l_2}{m_2} L - M \right) W_{\ell_1,\ell_2, M}^{XX} \phi_{LM},
\]

with the fields \( X_{\ell,m}, Z_{m',\ell'} \in \{T_{\ell,m}, E_{\ell,m}, B_{\ell,m}\} \). The covariance response functions \( W_{\ell_1,\ell_2, M}^{XX} \) for the possible field combinations can be found in Okamoto & Hu (2003). They are linear in the (lensed) CMB power spectra, and are only non-zero if \( l_1 + l_2 + L \) is even for even-parity combinations, e.g. \( TT \) and \( TE \), and \( l_1 + l_2 + L \) is odd for the odd-parity combinations \( TB \) and \( EB \). Furthermore, the \( W_{\ell_1,\ell_2, M}^{XX} \) are real for the even-parity combinations and imaginary for the odd-parity. We use the off-diagonal covariance to estimate the lensing potential; after the inverse-variance-filtered sky maps are generated, pairs of fields \( \tilde{X}_{\ell,m}, \tilde{Z}_{m',\ell'} \in \{T_{\ell,m}, \tilde{E}_{\ell,m}, \tilde{B}_{\ell,m}\} \) are fed into quadratic estimators to estimate the lensing potential. A general quadratic estimator is given as a function of input fields \( \tilde{X} \) and \( \tilde{Z} \) as

\[
\bar{s}_{LM}[\tilde{X}, \tilde{Z}] = \frac{(-1)^M}{2} \sum_{\ell,m,\ell',m'} \left( \frac{l_1}{m_1} \frac{l_2}{m_2} L \right) \times W_{\ell_1,\ell_2, M}^{XX} \tilde{X}_{\ell,m} \tilde{Z}_{m',\ell'},
\]

where \( W_{\ell_1,\ell_2, M}^{XX} \) is a set of weight functions that define the estimator \( x \), along with the input inverse-variance filtered multipoles \( \tilde{X}, \tilde{Z} \) for a particular map or pair of maps. The choice of \( \tilde{T}, \tilde{E}, \) or \( \tilde{B} \) for \( \tilde{X} \) and \( \tilde{Z} \) depends on the specific estimator \( x \). In most of what follows we shall drop the \( [\tilde{X}, \tilde{Z}] \) unless it is necessary to avoid ambiguity.

Optimal lensing estimators use a matched filter for the lensing-induced covariance of Eq. (A.13). We denote these as \( W_{XZ}^{XX} \) for the fields \( X \) and \( Z \). The temperature-only optimal estimator, for example, uses \( W^{TT} = W^{TT}_{\text{fid}} \), where the “fid” subscript indicates that the lensing weight function is evaluated for the fiducial cosmological model given at the end of Sect. 2. For the odd-parity combinations, the matched filter is minus the covariance response function, e.g. \( W^{TB} = -W^{TB}_{\text{fid}} \). Neglecting the lensing of any primordial B-mode signal (certainly a valid approximation for scales \( \ell_B > 100 \) with Planck polarization sensitivity), generally there are eight possible lensing estimators based on the possible combinations of \( T, E, \) and \( B \). (This reduces to five estimators for our baseline analysis that uses a foreground-cleaned map of the CMB in temperature and polarization, and for which there is no distinction between \( TE \) and \( ET \), for example.)

The quadratic estimators defined by Eq. (A.14) will receive contributions from non-lensing sources such as masking, beam asymmetry, and inhomogeneity of the instrumental noise. We determine and correct for this bias by averaging the reconstruction \( \bar{s}_{LM} \) over Monte Carlo simulations that include these effects, thereby estimating a “mean-field” \( S_{LM}^{MF} \) that we then subtract.

The response of a quadratic estimator to the covariance given in Eq. (A.13) (averaged over CMB realizations with a single mode \( \phi_{LM} \) of the lensing potential held fixed) is, in the diagonal approximation of Eq. (A.9),

\[
\langle \bar{s}_{LM} \rangle = K_{L}^{\text{rad}} \phi_{LM},
\]

where

\[
K_{L}^{\text{rad}} = \frac{1}{2(L + 1)} \sum_{\ell_1,\ell_2} W_{\ell_1,\ell_2, M}^{XX} F_{\ell_1}^{\tilde{T}} F_{\ell_2}^{\tilde{Z}}.
\]
As a baseline, all of the results in this paper are obtained using the fiducial model at the end of Sect. 2 to calculate the $\mathcal{W}^{N2}$ that appear in the estimator response. However, for cosmological parameter sampling with the lensing likelihood the normalization is calculated using the appropriate CMB power spectra as the parameter space is explored. This is discussed further in Appendix C.

Putting the above together, we form estimates of the lensing potential as

$$\hat{\phi}_{LM} = \frac{1}{R_L^{\phi}} \left( \hat{x}_{LM} - \bar{x}_{LM}^{MF} \right).$$  \hfill (A.17)

We can also sum the individual estimators into a combined minimum-variance estimator (MV) as

$$\hat{\phi}_{LM}^{MV} = \frac{\sum \hat{\phi}_{LM} R_L^{\phi}}{\sum R_L^{\phi}}.$$  \hfill (A.18)

where the sum is taken over the eight lensing estimators ($TT$, $EE$, $EE_T$, $EE_B$, $EB$, $ET$, $ET_B$, and $BE$).

Finally, we note an important implementation detail. Naively, estimating an estimator of the form in Eq. (A.14) requires $O(\ell_m^4 \ell_{\text{max}}^2)$ operations, which would be very computationally expensive at Planck resolution. Fortunately, all of the weight functions used in this paper can be rewritten as a sum of separable terms with the following form

$$W_{l_1, l_2} = \sum_j (-1)^{\ell_j} \frac{(2\ell_1 + 1)(2\ell_2 + 1)(2L + 1)}{4\pi} \left( \frac{\ell_1}{s_{l_1}} - \frac{\ell_2}{s_{l_2}} \right) \frac{L}{s_{l_2}^{l}} w_{l_1 l_2}^{x_j} x_{l_1 l_2}^{x_j}. \hfill (A.19)$$

This separability leads to estimators that can be evaluated using $O(\ell_{\text{max}}^4 \ell_{\text{max}})$ operations in position space as

$$\bar{x}_{LM} = \frac{1}{2} \int d\hat{n} \sum_{l_1, l_2} Y_{LM}(\hat{n}) w_{l_1 l_2}^{x_j} \left[ \sum_{j} Y_{l_1 l_2}(\hat{n}) w_{l_1 l_2}^{x_j} X_{l_1 l_2}^{x_j} \right]. \hfill (A.20)$$

The separability of the weight functions also allows the response functions to be calculated quickly analytically, using the position-space approach given in Dvorkin & Smith (2009) they can be evaluated in $O(\ell_{\text{max}}^4)$ operations.\footnote{A sample implementation of these weights can be found at http://github.com/dhanson/quicklens in the file quicklens/qest/lens.py}

A.3. Power spectrum estimation

Ensemble averaging over realizations of both the CMB lensing potential and the primary temperature and polarization fluctuations, the CMB becomes non-Gaussian. At first order in the lensing potential power spectrum $C_L^{\phi \phi}$ this non-Gaussianity is manifest as a connected 4-point function (Hu 2004)

$$X_{l_1 m_1} Z_{l_2 m_2} C_{l_3 m_3} D_{l_4 m_4} c = \sum_{LM} \left( \ell_1 \ell_2 L \ell_3 \ell_4 L \right) \left( m_1 m_2 M m_3 m_4 M \right) \times (-1)^{M+2} C_{L}^{\phi \phi} W_{l_1 l_2}^{N2} W_{l_3 l_4}^{CD} + 2 \text{ perms.} \hfill (A.21)$$

As might be expected given the presence of the covariance response functions $\mathcal{W}$ in this expression, estimators to extract $C_L^{\phi \phi}$ from the trispectrum can be written using the quadratic estimators above as building blocks. We write the cross-spectrum of two quadratic estimators explicitly as

$$C_{L_{xy}}^{\phi \phi} \equiv \frac{\ell_{xy}^{-1}}{2L+1} \sum M \phi_{LM}[\hat{X}, \hat{Z}] \phi_{LM}^{*}[\hat{C}, \hat{D}], \hfill (A.22)$$

where $f_{\text{sky}} = \sum p / N_{\text{pix}}$ is the unmasked sky fraction, with $M_p$ the mask map used in Eq. (A.7). We form estimates of the lensing potential power spectrum (based on estimators $x$ and $y$) as

$$\hat{C}_{L_{xy}}^{\phi \phi} = C_{L_{xy}}^{\phi \phi} - \Delta C_{L_{xy}}^{\phi \phi} \bigg|_{N0} - \Delta C_{L_{xy}}^{\phi \phi} \bigg|_{N1} - \Delta C_{L_{xy}}^{\phi \phi} \bigg|_{MC} - \Delta C_{L_{xy}}^{\phi \phi} \bigg|_{PS}. \hfill (A.23)$$

There are several correction terms here, which are discussed in more detail below. We also plot them for the MV reconstruction based on the Planck 2015 SMICA maps in Fig. A.1.

The $\Delta C_{L_{xy}}^{\phi \phi} \bigg|_{N0}$ term represents the disconnected contribution to the 4-point function, which would be non-zero even in the absence of lensing. We estimate this term by replacing some of the data fields with those from two sets of independent simulations (labelled MC1 and MC2), and then average over realizations as

$$\Delta C_{L_{xy}}^{\phi \phi} \bigg|_{N0} = -C_{L_{xy}}^{\phi \phi} \bigg[ \hat{X}_{MC1}, \hat{Z}_{MC2}, \hat{C}_{MC1}, \hat{D}_{MC1} \bigg]$$

$$+ C_{L_{xy}}^{\phi \phi} \bigg[ \hat{X}_{MC1}, \hat{Z}, \hat{C}_{MC1}, \hat{D} \bigg]$$

$$+ C_{L_{xy}}^{\phi \phi} \bigg[ \hat{X}, \hat{Z}_{MC1}, \hat{C}_{MC1}, \hat{D} \bigg]$$

$$- C_{L_{xy}}^{\phi \phi} \bigg[ \hat{X}, \hat{Z}_{MC1}, \hat{C}_{MC2}, \hat{D}_{MC1} \bigg]. \hfill (A.24)$$

This method for determining the disconnected bias can be derived from the Edgeworth expansion of the lensed CMB (Regan et al. 2010; neglecting the $T\phi$ and $E\phi$ correlations). If the simulations have a covariance differing from reality by a linear (in

Fig. A.1. Lens reconstruction bias terms for the MV lensing estimator applied to the Planck 2015 SMICA maps. Dashed lines indicate regions where the bias term is negative. The fiducial $\Lambda$CDM theory power spectrum $C_L^{\phi \phi} \bigg|_{PS}$ is plotted as the black solid line.
general, anisotropic) correction, this form of the N0 estimator using a mixture of data and simulations is insensitive to the correction at linear order (Namikawa et al. 2013). The largest difference we see between simulations and data is about 5% in power before correction, so the quadratic error should be at most about $(0.05)^2 = 0.0025$, corresponding to an error of less than 0.4σ for aggressive bins at $L \leq 400$ rising to around 1.3σ at $L = 2000$. Since we also match the data calibration, and add isotropic foreground and noise power to the simulations to match the observed amplitude and shape of the power spectrum of the data, the isotropic part of the error is actually close to zero. A residual error only arises at second order due to our inability to capture any statistical anisotropy in the missing foreground and noise power, and hence should be much smaller than the above estimates. This makes the N0 subtraction rather robust, despite it being substantially larger than the signal on small scales. Furthermore, the statistical anisotropy in the missing foreground and noise power spectrum estimates between different multipoles (Hansen et al. 2011) and with the measured CMB power spectra (Schmittfull et al. 2013).

The $\Delta C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4}$ term corrects for the non-Gaussian secondary contractions (the other permutations in Eq. A.21) due to lensing, as discussed in Kesden et al. (2003). We evaluate it using the flat-sky approximation in 2D Fourier space as

$$\Delta C^{\phi \phi}_{\ell, \nu|N_1} = \frac{1}{R^2_{\ell, f} R^2_{\nu, f}} \int \frac{d^2 l_1}{(2\pi)^2} \int \frac{d^2 l_2}{(2\pi)^2} \frac{d^2 l_3}{(2\pi)^2} \frac{d^2 l_4}{(2\pi)^2} \times \left[ C^{\phi \phi}_{\ell_1 \ell_2} W_{\ell_1 \ell_2} (-l_1, l_1') W_{\ell_2 \ell_2'} (-l_2, l_2') + C^{\phi \phi}_{\ell_1 \ell_3} W_{\ell_1 \ell_3} (-l_1, l_3') W_{\ell_3 \ell_3'} (-l_3, l_3') \right] \times \frac{F_{l_1} F_{l_2} F_{l_3} F_{l_4}}{2^4} \times W^n(-l_1, l_1) W^n(-l_2, l_2),$$

where $l_1 + l_2 = l_3 + l_4 = L$. Here the weight functions are flat-sky lensing weight functions; they can be found in Hu & Okamoto (2002).

We characterize any differences between the average lensing power obtained on simulations with the input power. This is determined by

$$\Delta C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4} \equiv \left( C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4}[X_{\ell_1}, Z_{\ell_1}, X_{\ell_2}, Z_{\ell_2}] \right) - C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4}[X_{\ell_1}, Z_{\ell_1}, X_{\ell_2}, Z_{\ell_2}] - C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4}[X_{\ell_1}, Z_{\ell_1}, X_{\ell_2}, Z_{\ell_2}] - C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4}[X_{\ell_1}, Z_{\ell_1}, X_{\ell_2}, Z_{\ell_2}] - \Delta C^{\phi \phi}_{\ell, \nu|N_1} = \left( C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4}[X_{\ell_1}, Z_{\ell_1}, X_{\ell_2}, Z_{\ell_2}] \right)_{MC2} \times \left( C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4}[X_{\ell_1}, Z_{\ell_1}, X_{\ell_2}, Z_{\ell_2}] \right)_{MC1}. \quad (A.26)$$

The MC correction is included to account for possible issues in the estimation of the normalization, mixing of power between power spectrum bins, and errors in the calculation of the N1 bias. Although in principle several of these effects should be included as multiplicative rather than additive corrections, we find that the MC correction is generally small enough (less than 10% of the fiducial lensing power spectrum) that it does not matter at our experimental sensitivity whether we apply the correction additively or multiplicatively.

Finally, we make a correction $\Delta C^{\phi \phi}_{\ell, \nu|PS}$ for the non-Gaussianity due to the shot-noise of unresolved point sources in temperature. This is given by

$$\Delta C^{\phi \phi}_{\ell, \nu|PS} = \tilde{S}^2 \frac{\tilde{R}^{X_{\ell_1} X_{\ell_2}} \tilde{R}^{Z_{\ell_1} Z_{\ell_2}}}{\tilde{R}^{X_{\ell_1}} \tilde{R}^{Z_{\ell_2}}} \left[ \sum_{m} \tilde{X}_{\ell_1} \tilde{Z}_{\ell_2} \right]. \quad (A.27)$$

where $\tilde{S}^2$ is an estimate of the shot-noise amplitude and $S^2$ denotes the weight function (non-zero only for $TT$) defined by

$$W^{S^2}_{\ell_1 \ell_2 L} = \frac{(2\ell_1 + 1)(2\ell_2 + 1)(2L + 1)}{4\pi} \left[ \begin{array}{c} \ell_1 \\ \ell_2 \\ L \end{array} \right]. \quad (A.28)$$

The $R^{S^2}$ is given by Eq. (A.16) but with $W^{\phi \phi}_{\ell_1 \ell_2 L}$ replaced by $W^{S^2}_{\ell_1 \ell_2 L}$. Following Osborne et al. (2014), the estimate for the shot-noise trispectrum amplitude is obtained in an analogous way to Eq. (A.23), but using the quadratic estimator formed with the point-source weight function above rather than the lensing weight function. The overall source amplitude is then estimated as

$$\hat{S}^2 = \frac{\sum_{\ell}(2L + 1) \left( R_{\ell}^{S^2} \right)^2 C^S_{\ell} \tilde{C}^S_{\ell} \tilde{C}^S_{\ell}}{\sum_{\ell}(2L + 1) \left( R_{\ell}^{S^2} \right)^2 C^S_{\ell} C^S_{\ell}}, \quad (A.29)$$

where the sums are taken over $10 \leq L \leq 2048$. Note that $R^{S^2}$ is given by Eq. (A.16) but with both $W^{\phi \phi}_{\ell_1 \ell_2 L}$ and $W^{\phi \phi}_{\ell_1 \ell_2 L}$ replaced by $W^{S^2}_{\ell_1 \ell_2 L}$.

We use Eq. (A.23) to estimate the lensing potential power spectrum, which we then bin into bandpowers for plotting as well as constructing a lensing likelihood. We use Monte Carlo simulations of the power spectrum estimation to construct error bars for these bandpowers. There is one technical point here, which is that the realization-dependent $\Delta C^{\phi \phi}_{\ell, \nu|PS}$ correction term can be cumbersome to compute. Instead, for the purpose of determining error bars, we use a semi-analytical approximation to this term given by

$$\Delta C^{\phi \phi}_{\ell, \nu|N_0, \mathrm{analytic}} \left[ \tilde{X}, \tilde{Z}, \tilde{C}, \tilde{D} \right] = \frac{1}{R_{\ell}^{\phi \phi} R_{\nu}^{\phi \phi}} \frac{1}{4(2L + 1)} \times \sum_{\ell_1 \ell_2} W^\phi_{\ell_1 \ell_2 L} \left( C^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4} \tilde{C}^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4} \right)^2 \left( X^{\phi \phi}_{\ell_1 \ell_2} \tilde{X}^{\phi \phi}_{\ell_1 \ell_2} \tilde{C}^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4} \tilde{C}^{\phi \phi}_{\ell_1 \ell_2 \ell_3 \ell_4} \right), \quad (A.30)$$

where the empirical power spectra of the filtered fields are

$$\hat{C}_{\ell}^{XX} = \frac{f_{\ell, \text{sky}}}{2L + 1} \sum_{m} \tilde{X}_{\ell_1} \tilde{Z}_{\ell_2}. \quad (A.31)$$

Appendix B: Bias-hardened estimators

The quadratic lensing estimators of Appendix A.2 can be biased by non-lensing sources of statistical anisotropy such as beam asymmetry, noise inhomogeneity, and masking. In our fiducial analyses, we estimate these mean-field biases using Monte Carlo simulations. However, this approach can be sensitive to the fidelity of the simulations, particularly for the reconstruction on large angular scales where the mean-field biases can be orders of magnitude above the lensing signals of interest. As a cross-check, we also take a more data-dependent approach to correcting these biases. As with lensing, most sources of statistical anisotropy in a map can be associated with a covariance response function $W_{\ell_1 \ell_2 L}$, and a field $Z_{LM}$ that describes the spatial dependence of the anisotropy. It is possible to construct quadratic estimators for $Z_{LM}$, and then calculate and subtract the resulting lensing bias. This procedure amounts to constructing special “bias-hardened” estimators (Namikawa et al. 2013). Given
a source of bias, we construct a bias-hardened estimator with weight function
\[ W_{\ell_{1} \ell_{2}; \ell} = W_{\ell_{1} \ell_{2}} - \frac{\mathcal{R}_{\ell_{1}}^{\text{c}}}{\mathcal{R}_{\ell_{3}}^{\text{c}}} W_{\ell_{3} \ell_{2}; \ell}, \tag{B.1} \]
where \( \mathcal{R}_{\ell_{1}}^{\text{c}} \) and \( \mathcal{R}_{\ell_{3}}^{\text{c}} \) are generalizations of Eq. (A.16), and \( W_{\ell_{1} \ell_{2}} \) is the matched filter for \( W_{\ell_{1} \ell_{2}; \ell} \). This method is potentially less susceptible to errors in the mean-field estimate (due to inaccuracies in the simulations used to obtain it), although it has several caveats that are discussed in more detail in Namikawa et al. (2013).

Appendix C: Power spectrum likelihood

Here, we provide more detail on the components of our lensing likelihood (given in Eq. (3), and copied below for convenience)
\[ -2 \log L_{\theta} = B_{C}^{i}(C_{\ell}^{\phi \phi} - C_{\ell}^{\phi \phi; \text{fid}})[\Sigma^{-1}]^{-1/2} B_{C}^{j}(C_{\ell}^{\phi \phi} - C_{\ell}^{\phi \phi; \text{fid}}). \tag{C.1} \]
The bandpower binning function \( B_{C}^{i} \) is given by
\[ B_{C}^{i} = \frac{C_{\ell}^{\phi \phi; \text{fid}} V_{L}^{-1}}{\sum_{L} \left( C_{\ell}^{\phi \phi; \text{fid}} \right)^{2}} V_{L}, \tag{C.2} \]
where \( V_{L} \) is an approximation to the variance of the power spectrum estimate. For a given cross-spectrum between estimators \( x \) and \( y \) we use
\[ V_{L}^{-1} = \frac{2L+1}{2\ell_{\text{sky}}} \mathcal{R}_{\ell_{x}}^{\phi} \mathcal{R}_{\ell_{y}}^{\phi}, \tag{C.3} \]
which produces minimum-variance estimates of the lensing amplitude for each bin in the limit that the sample variance of the lenses is negligible (a reasonable approximation for Planck, which is noise dominated on any individual mode).

The dependence of the likelihood on cosmological parameters enters through the “theory” spectrum \( C_{\ell}^{\phi \phi; \text{fid}} \). This is the expected value of the estimated spectrum at each point in parameter space. It depends on the cosmological parameters \( \theta \) in several ways:

- directly, through the theory spectrum \( C_{\ell}^{\phi \phi; \text{fid}} \); and
- indirectly (but linearly) on \( C_{\ell}^{\phi \phi; \text{fid}} \), via the theory-dependence of \( N^{(1)} \) (Eq. (A.25)); and
- indirectly, and non-linearly, on the CMB power spectra \( C_{TT}, C_{EE}, \) and \( C_{\ell}^{\phi \phi} \) through the estimator normalization and \( N^{(1)} \) (from the theory-dependent covariance \( \mathcal{W} \) of Eq. (A.13)).

We neglect any other theory dependence, for example in the MC correction that is hard to quantify. In detail, for a given set of cosmological parameters \( \theta \), we should calculate
\[ C_{\ell}^{\phi \phi; \text{fid}} = \frac{(\mathcal{R}_{\ell}^{\phi}, \mathcal{R}_{\ell}^{\phi})}{(\mathcal{R}_{\ell}^{\phi}, \mathcal{R}_{\ell}^{\phi})_{\text{fid}}} \left[ C_{\ell}^{\phi \phi; \text{fid}} - \Delta C_{\ell}^{\phi \phi; \text{fid}} \right] + \Delta C_{\ell}^{\phi \phi; \text{fid}}, \tag{C.4} \]
where the “\( \mathcal{W} \)” component of the response functions \( \mathcal{R} \) in the numerator, as well as the \( N^{(1)} \) term, are calculated using the CMB power spectra for parameters \( \theta \).

Fully recalculating everything for each point in parameter space is prohibitively slow, but for small deviations from the fiducial model we can use a linearized approximation. The dependence on \( C_{\ell}^{\phi \phi} \) is already linear, and expanding the CMB power spectrum dependence to linear order about the fiducial model we can write
\[ C_{\ell}^{\phi \phi; \text{fid}} = C_{\ell}^{\phi \phi} + \frac{\text{dln}(\mathcal{R}_{\ell}^{\phi})}{\text{d}C_{\ell}^{\phi}} \left( C_{\ell}^{\phi \phi} - C_{\ell}^{\phi \phi; \text{fid}} \right) + \frac{\text{dCN}_{\ell}^{\phi \phi; \text{fid}}}{\text{d}C_{\ell}^{\phi}} \left( C_{\ell}^{\phi \phi} - C_{\ell}^{\phi \phi; \text{fid}} \right), \tag{C.5} \]
where \( j \) sums over the various CMB power spectra. The matrix \( N_{\ell}^{(1)} \) (from Eq. (A.25)) gives the exact linear dependence of \( N^{(1)} \) on the lensing potential for fixed CMB power spectra, and can be pre-computed along with the other derivative matrices for the fiducial model. The binned likelihood of Eq. (C.1) then depends on a binned “theory” power spectrum given in the linear approximation by
\[ B_{C}^{i} C_{\ell}^{\phi \phi; \text{fid}} = B_{C}^{i} C_{\ell}^{\phi \phi} + M_{\ell}^{i; \ell} \left( C_{\ell}^{\phi \phi} - C_{\ell}^{\phi \phi; \text{fid}} \right), \tag{C.6} \]
where now \( a \) sums over both the \( C_{\ell}^{\phi \phi} \) and CMB power spectra terms, and \( M_{\ell}^{i; \ell} \) can be pre-computed in the fiducial model.

The Planck lensing likelihood consists of bandpower estimates \( B_{C}^{i} C_{\ell}^{\phi \phi} \) and the covariance \( \Sigma_{ij} \), a set of binning functions \( B_{C}^{i} \), a set of linear-correction kernels \( M_{\ell}^{i; \ell} \), and the band values of the linear-correction kernel applied to the fiducial spectrum \( M_{\ell}^{i; \ell} C_{\ell}^{\phi \phi; \text{fid}} \). Note that since the lensing estimator is independent of the CMB calibration, for a theory model \( C_{\ell}^{\phi \phi} \) and map calibration parameter \( y_{\text{cal}} \), the theory CMB power spectra that are used in Eq. (C.6) are those matched to the data, i.e. \( C_{\ell}^{\phi \phi} = C_{\ell}^{\phi \phi; y_{\text{cal}}} \).

When evaluating “lensing-only” likelihoods, where the CMB power spectra are only very weakly indirectly constrained, we fix the \( C_{\ell}^{\phi \phi} \) used in the linear correction to a best-fit to the full data. When calculating joint lensing and CMB results, the correction depends on the CMB power spectra at each point in model space.

Appendix D: CIB model

The cosmic infrared background (CIB) is a diffuse sky signal which begins to dominate over the CMB at frequencies \( \nu \gtrsim 300 \) GHz. It is generated by dust that is heated by UV light from young stars and then re-radiates thermally in the infrared. The CIB contains approximately half of the total extragalactic stellar flux, and has excellent redshift overlap with the CMB lensing potential (Song et al. 2003). In Sect. 3.2 we use the CIB as a tracer of the lensing potential to estimate the lensing-induced B-mode signal. This requires a model for the cross-correlation between the CIB and the lensing potential. We calculate this in the Limber approximation as (Limber 1954)
\[ C_{L}^{\phi \phi; \text{CIB}} = \int \frac{\text{d}k}{\chi} K_{L}^{\phi \phi; \text{CIB}}(\chi) K_{L}^{\phi \phi}(\chi) P(k = L/\chi; \eta_{0} - \chi), \tag{D.1} \]
where \( P(k; \eta) \) is the (non-linear) matter power spectrum at comoving wavenumber \( k \) and conformal time \( \eta \). The lensing kernel in flat models is given by
\[ K_{L}^{\phi \phi}(\chi) = \frac{3\Omega_{m} H_{0}^{2}}{\chi(\chi^{2} + L^{2})} \left( \frac{\chi^{2} - \chi}{\chi L} \right), \tag{D.2} \]
where $d(\chi)$ is the scale factor at conformal time $\eta_0 - \chi$ and $z_c$ denotes the last scattering surface at $z \approx 1100$. To model the CIB, we use the simple SSED model of Hall et al. (2010), which has a kernel given by

$$K^\text{CIB}_L(\chi) = b_c \frac{\chi^2}{(1+z)^2} \exp\left(-\frac{(z-z_c)^2}{2\sigma_z^2}\right) f_r(1+z),$$

where $b_c$ is an overall normalization and $z_c = \sigma_z = 2$ describe the redshift distribution of the CIB intensity. The redshift $z$ is evaluated at conformal time $\eta_0 - \chi$. The spectral energy distribution (SED) of a typical CIB source is described by $f_r$. It is modelled as a modified blackbody with temperature $T = 34$ K, and spectral index $\beta = 2$ up to a threshold $\nu' = 4955$ GHz, where it transitions to a power-law decay with index $\alpha = 2$:

$$f_r = \begin{cases} \exp\left(\frac{\nu}{\nu'} - 1\right)^{-1} & (\nu \leq \nu'), \\ \exp\left(\frac{\nu'}{\nu} - 1\right)^{-1} \nu'^{\beta+3} \left(\frac{\nu}{\nu'}\right)^{-\alpha} & (\nu > \nu') \end{cases}.$$  

The precise value of $\nu'$ is chosen to match smoothly the slope of $f_r$ on both sides of the transition frequency. In this model, the modified blackbody component of the SED is generally what we observe; even at an observation frequency of 1200 GHz (i.e. 250 µm), the power-law transition does not occur until $z \approx 4$. Our theory curve is plotted in Fig. D.1, as well as the measured cross-correlation of the Planck 2015 545 GHz map with the MV lensing potential estimate (over 40% of the sky). In addition we have plotted the bandpowers from the rigorous analysis presented in the Planck 2013 data release (Planck Collaboration XVIII 2014). We have set the normalization $b_c$ for the theory curve such that the overall amplitude is consistent with both measurements.

**Appendix E: Parameter dependence of $C_L^\Phi$ in $\Lambda$CDM models**

In this appendix we discuss the dependence of the lensing potential power spectrum at multipoles $L \gtrsim 100$ on the parameters of $\Lambda$CDM models, providing results that are used in the discussion of the “lensing-only” constraints in Sect. 3.5. Since these constraints adopt a prior on $n_s$ that is rather tighter than the resulting posterior widths for $\Omega_m h^2$ and $h$, we ignore dependencies on $n_s$ throughout this appendix. We caution the reader that for joint constraints, or fits to the CMB power spectra alone, fractional variations in $n_s$ can be comparable to those in $\Omega_m h^2$ and $h$, so the dependence of $C_L^\Phi$ on $n_s$ cannot be ignored. The dependence of $C_L^\Phi$ on the matter density was recently discussed in Pan et al. (2014); we follow their discussion closely.

For multipoles $L \gtrsim 20$, the angular power spectrum of the lensing potential can be Limber-approximated as

$$L^3 C_L^\Phi = 4 \int_0^\chi \frac{d\chi}{\chi} (k^2 P_\Phi(k)) (k/L|\eta_0 - \chi|) (1 - \chi/\chi_c)^2$$

in spatially-flat models. Here, $P_\Phi(k;\eta)$ is the equal-time dimensional power spectrum of the Weyl potential $\Psi$. Well after baryon decoupling, we have

$$k^3 P_\Phi(k;\eta) \propto k^3 P_{B}(T^2(k) g^2(\nu)),$$

where $P_B(k)$ is the dimensionless power spectrum of the primordial curvature perturbation $\mathcal{R}$, $T(k)$ is the transfer function, and $g(\nu)$ is the growth function for the Weyl potential at scale factor $a$, normalized to unity at high redshift. In $\Lambda$CDM,

$$g(a) = \frac{5\Omega_m E(a)}{2a} \left[ \int_0^\infty \frac{da'}{a'(E(a'))} \right].$$

where $\Omega_m$ is the matter fraction at the present and $E(a) \equiv H(a)/H_0$. For $\Omega_m \approx 0.3$, the current value of the growth function is $0.78$ and $d\ln G/d\ln \Omega_m = 0.23$. Close to $\Omega_m = 0.3$, we can approximate the cosmology dependence of the growth function at redshift zero as $g(a) = 0.78(\Omega_m/0.3)^{0.23}$. For $L > 100$, the lensing potential mostly probes early enough times that $g(a) \approx 1$ is a good approximation (Pan et al. 2014). However, if we wish to express the amplitude of $P_\Phi(k)$ at early times in terms of $\sigma_8$ (at the present), we require $g(a) = 1$.

The transfer function is normalized to unity for $k \ll k_{eq}$, where $k_{eq} \sim 10^{-2}$ Mpc$^{-1}$ the inverse of the comoving Hubble parameter at matter-radiation equality. For $k \gg k_{eq}$ we have $T(k) \sim (k/k_{eq})^{-2} \ln(k/k_{eq})$. Keeping the radiation energy density fixed today, $k_{eq} \propto \omega_m$, where $\omega_m \equiv \Omega_m h^2$. As well as $k_{eq}$, the other scales that enter the transfer function are the sound horizon at the drag epoch (when baryons dynamically decouple) $r_{drag}$ and the diffusion scale $k_D$. The broad-band shape is controlled by $k_{eq}$, while $r_{drag}$ determines the scale below which baryon suppression effects are important (see e.g. Eisenstein & Hu 1998). The non-zero baryon fraction $\omega_B/\omega_m$ tends to suppress the transfer function on small scales since (essentially) unclustered baryons impede the growth of the CDM density perturbation up to the time after decoupling when baryon infall into the CDM potential wells is complete. The further effect of baryons is to imprint baryon acoustic oscillations in the transfer function for $k_{drag} > \pi$ and $k \lesssim k_b$. Baryon acoustic oscillations have little impact on $C_L^\Phi$ since they are washed out by the line-of-sight integration. We therefore ignore them here and, where an explicit form for the transfer function is required, for illustrative purposes we use the “no-wiggle” fit of Eisenstein & Hu (1998).

For the range of parameters of interest, we can assume that the

15 We ignore the small effects of non-zero neutrino mass in the base $\Lambda$CDM model, so that the linear growth of structure is scale-independent well after baryon decoupling.
physical baryon density $\omega_b$ is fixed, so the transfer function can be thought of as a function of $k/k_{eq}, kr_{drag}$ and $\omega_m$. Ignoring departures of the growth function from unity and approximating $P_{\gamma}(k)$ as scale-invariant with amplitude $A_s$, we can write

$$L^4 C^b_{L^*} \propto A_s \ell_{eq} \int_0^1 dx \left( \frac{L}{\ell_{eq}} \right)^2 \left( \frac{L}{X_s} \right)^2 (1 - x)^2,$$  \hspace{1cm} (E.4)

where $x = k/k_{eq}$, and $\ell_{eq} \equiv k_{eq} X_s \approx 140$. The $A_s \ell_{eq}$ prefactor has a simple physical interpretation: CMB photons are typically deflected by $\ell_{eq}$ lenses of size $k_{eq}^{-1}$, and each mean-squared deflection is proportional to $A_s$, giving a mean-squared deflection proportional to $A_s \ell_{eq}$. Given the $k$-dependencies of the transfer function through $k/k_{eq}$ and $kr_{drag}$, and its direct parameter-dependence on $\omega_m$, we see that in $\Lambda$CDM, the lensing power spectrum at $L \geq 100$ is (mostly) determined by only $A_s, L/\ell_{eq}, L/\ell_{drag}$, and $\omega_m$. Here, the multipole $\ell_{drag} \equiv \chi_x/r_{drag}$ characterizes the BAO scale at last-scattering. The integral in Eq. (E.4) determines the shape of the lensing power spectrum. Defining

$$f = \frac{L}{\ell_{eq}}, \quad \frac{L}{\ell_{drag}}, \quad \omega_m,$$

the dependence on the cosmological parameters (e.g. $\ell_{eq}$) in the vicinity of the fiducial model is, at any $L$, approximately a product of power laws with exponents

$$n_{eq}^{\chi_x} \equiv -\partial \ln f_{\ell_{eq}}/\partial \ln (L/\ell_{eq}) \bigg|_{\text{fid}}, \quad (E.6)$$

$$n_{drag}^{\chi_x} \equiv -\partial \ln f_{\ell_{drag}}/\partial \ln (L/\ell_{drag}) \bigg|_{\text{fid}}, \quad (E.7)$$

$$n_{L^*}^{\chi_x} \equiv \partial \ln f_{\ell_{drag}}/\partial \ln \omega_m \bigg|_{\text{fid}}, \quad (E.8)$$

each evaluated for the fiducial cosmology. Note also that $n_{eq}^{\chi_x}$ and $n_{drag}^{\chi_x}$ determine the local power-law slope of $C^b_{L^*}$, with

$$n_{eq}^{\chi_x} + n_{drag}^{\chi_x} = -\partial \ln (L^4 C^b_{L^*})/\partial \ln L.$$  \hspace{1cm} (E.9)

On the angular scales where Eq. (E.4) is valid, only modes with $k > k_{eq}$ contribute. For these scales, $(k/k_{eq})^3 T^2(k)$ increases with $k_{eq}$ giving $n_{eq}^{\chi_x} > 0$ (and monotonic increase of $n_{eq}^{\chi_x}$ with $L$). Around $L \geq 200$, the scales that dominate the $S/N$ of a lensing amplitude measurement from Planck, $n_{eq}^{\chi_x} \approx 0.9$. This, and $A_s$, dominate the parameter dependence of $C^b_{L^*}$. The effect of baryon suppression is reduced as $\omega_m$ increases at fixed $\omega_b$, so $n_{eq}^{\chi_x} > 0$ and also rises monotonically with $L$. Around $L \leq 200$, we find $n_{eq}^{\chi_x} \approx 0.3$, which gives a small, but non-negligible, further parameter dependence to $C^b_{L^*}$. Finally, the dependence on $\ell_{drag}$ is very small, since $n_{drag}^{\chi_x} \leq 0.15$ (with peak value around $L \approx 100$, roughly equal to $\ell_{drag}$). Putting these pieces together, we find a dependence of $L^4 C^b_{L^*}$ on cosmological parameters scaling approximately as

$$L^4 C^b_{L^*} \propto A_s \ell_{eq} n_{eq}^{\chi_x} + n_{drag}^{\chi_x} \omega_m^m.$$  \hspace{1cm} (E.10)

for $L \geq 100$, where the dependence on $\ell_{drag}$ is very weak.

We can express the parameter dependence of $C^b_{L^*}$ in terms of the usual $\Lambda$CDM parameters. The comoving distance to last-scattering, $\chi_x$, is

$$\chi_x = \int_0^1 \frac{da}{a \sqrt{H(a)}},$$  \hspace{1cm} (E.11)

where $a_s$ is the scale factor at last-scattering (whose dependence on cosmology can be ignored here). At fixed radiation energy density today, the dependence on $\Omega_m$ and $h$ is approximately

$$\chi_x \propto \Omega_m^{-0.4} h^{-1},$$  \hspace{1cm} (E.12)

around $\Omega_m = 0.3$ and $h = 0.7$. This differs from the $(\Omega_m h^2)^{-1/2}$ of an Einstein-de Sitter universe due to the late-time effects of $A$. Using $\ell_{eq} \propto \Omega_m h$, we have

$$\ell_{eq} \propto \Omega_m^{-0.15} h^{-0.5}.$$  \hspace{1cm} (E.13)

At fixed baryon density, the sound horizon at the drag epoch scales like $r_{drag} \propto \omega_m^{-0.25}$ in the vicinity of the fiducial model, so that

$$\ell_{drag} \propto \Omega_m^{0.15} h^{-0.5}.$$  \hspace{1cm} (E.14)

Using these dependencies in Eq. (E.10), we find

$$L^4 C^b_{L^*} \propto A_s \left( \Omega_m^{0.6} h^{1.3} \right) \left( \Omega_m^{-0.15} h^{-0.5} \right),$$  \hspace{1cm} (E.15)

where $n_L \equiv n_{eq}^{\chi_x} - 0.5 \times n_{drag}^{\chi_x} + 2n_{L^*}^{\chi_x}$ is approximately 1.5 at $L = 200$. The remaining dependence on $\Omega_m$ is very weak, with $0.15 \times n_{drag}^{\chi_x} - 0.2 \times n_{L^*}^{\chi_x} \approx -0.05$ at $L = 200$. Equation (E.15) gives the main dependence of the lensing power spectrum on the cosmological parameters of $\Lambda$CDM models at multipoles $L \geq 100$.

### E.1. Normalization by $\sigma_8$

The dependence of the lensing power spectrum on $A_s$ can be eliminated in favour of $\sigma_8$. We use the Poisson equation to relate the dimensional matter power spectrum at $a = 1$ to that of the Weyl potential, giving

$$P_\gamma(k; a = 1) \propto \omega_m^{-2} k^4 P_\gamma(k; a = 1).$$  \hspace{1cm} (E.17)

It follows that the variance of the density contrast in spheres of radius $R$ is

$$\sigma_R^2 \propto \omega_m^{-2} \sigma_8^2 (a = 1) \int_0^R dk k^4 P_\gamma(k) T^2(k) W^2(kR) \propto A_s k_0^4 \omega_m^{-2} \sigma_8^2 (a = 1) \int_0^\infty dk k^3 T^2(kx_m) W^2(kx_m R),$$  \hspace{1cm} (E.18)

where we have assumed that $P_\gamma(k)$ is scale-invariant in passing to the second line. The window function

$$W(kR) \equiv \frac{3J_1(kR)}{kR} = \frac{3}{(kR)^2} (kR \cos kR - \sin kR),$$  \hspace{1cm} (E.19)

is the Fourier transform of a normalized spherical top-hat function of radius $R$; following convention, we take $R = 8h^{-1}$ Mpc.

The integral in Eq. (E.18) is a function of $k_0 R, k_0 r_{drag}$, and $\omega_m$ at fixed baryon density (with the latter two from the transfer function). Evaluating the logarithmic derivative with respect to these parameters at the fiducial values, using the “no-wiggle” transfer function from Eisenstein & Hu (1998), we find that the integral scales as $(k_0 R)^{-1} \omega_m^{-0.45}$, with only a very weak further dependence on $k_0 r_{drag}$, in the vicinity of the fiducial model. Finally, the parameter dependence of $\sigma_R^2$ is

$$\sigma_R^2 \propto A_s k_0^4 \omega_m^{-2} \sigma_8^2 (a = 1) (k_0 h^{-1})^{-1.4} \Omega_m^{-0.45} \propto A_s \Omega_m^{1.5} h^{-3.5}.\quad (E.20)$$

This parameter dependence agrees well with direct finite-differencing of $\sigma_8$ calculations with CAMB.

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A15, page 28 of 28

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